

Fermi GBM and Spectral Fitting GRBs

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University of
Alabama
in Huntsville



NASA
Marshall Space Flight
Center



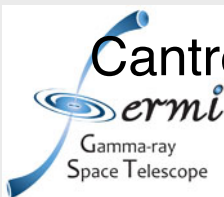
Max-Planck-Institut für
extraterrestrische Physik

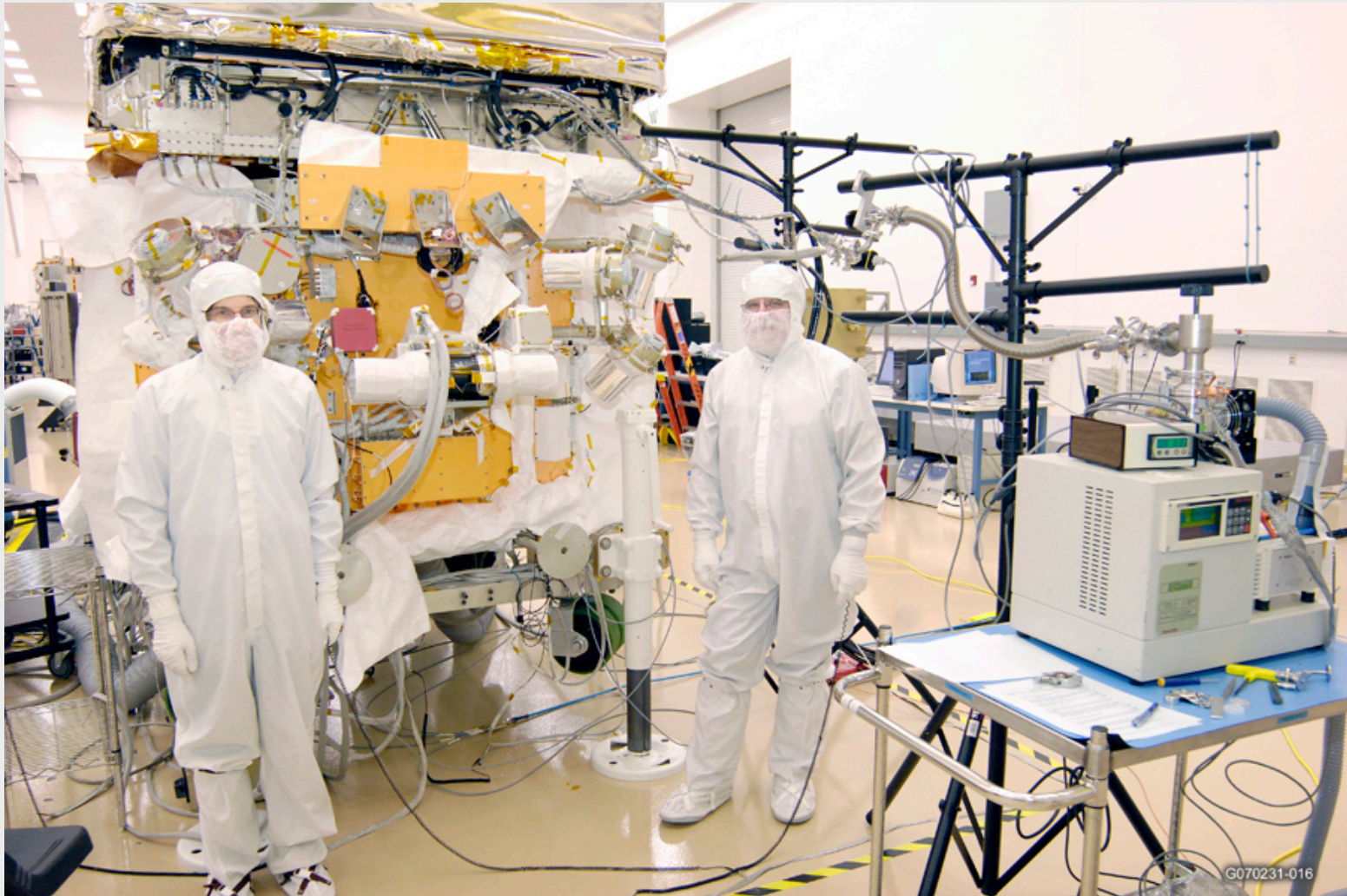


University
College
Dublin

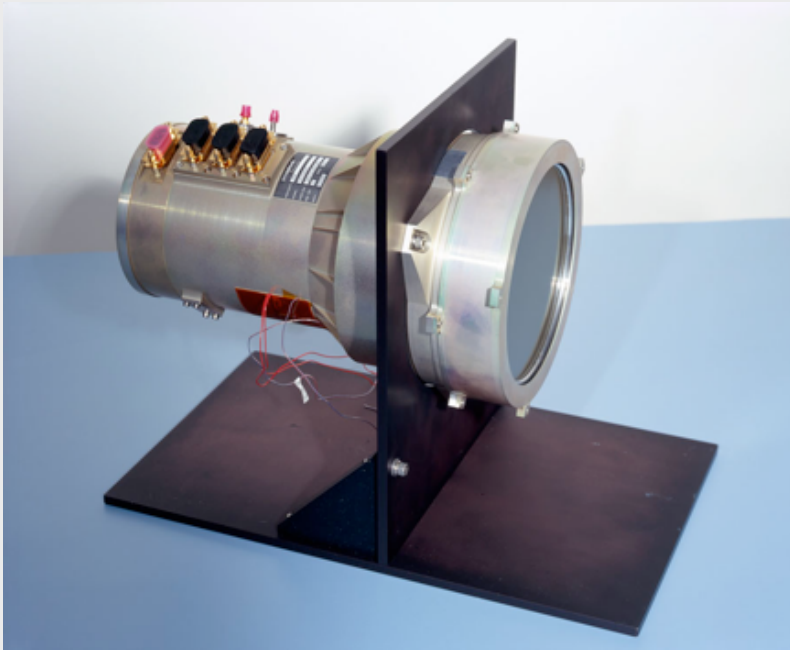
The current GBM Team: Narayana Bhat, Michael Briggs, Michael Burgess, Vandiver Chaplin, Bill Cleveland, Valerie Connaughton, Roland Diehl, Steve Elrod, Mark Finger, Jerry Fishman, Gerard Fitzgerald, Suzanne Foley, Lisa Gibby, Misty Giles, Adam Goldstein, Jochen Greiner, David Gruber, Alexander van der Horst, Andreas von Kienlin, Pete Jenke, Marc Kippen, Chryssa Kouveliotou, Emily Layden, Sheila McBreen, Sinead McGlynn, Chip Meegan, Bill Paciasas, Veronique Pelassa, Rob Preece, Arne Rau, Dave Tierney, Colleen Wilson-Hodge and Shaolin Xiong.

Others before launch: Giselher Lichti, Fred Berry, Ron
AI English, Fred Kroeger, ...

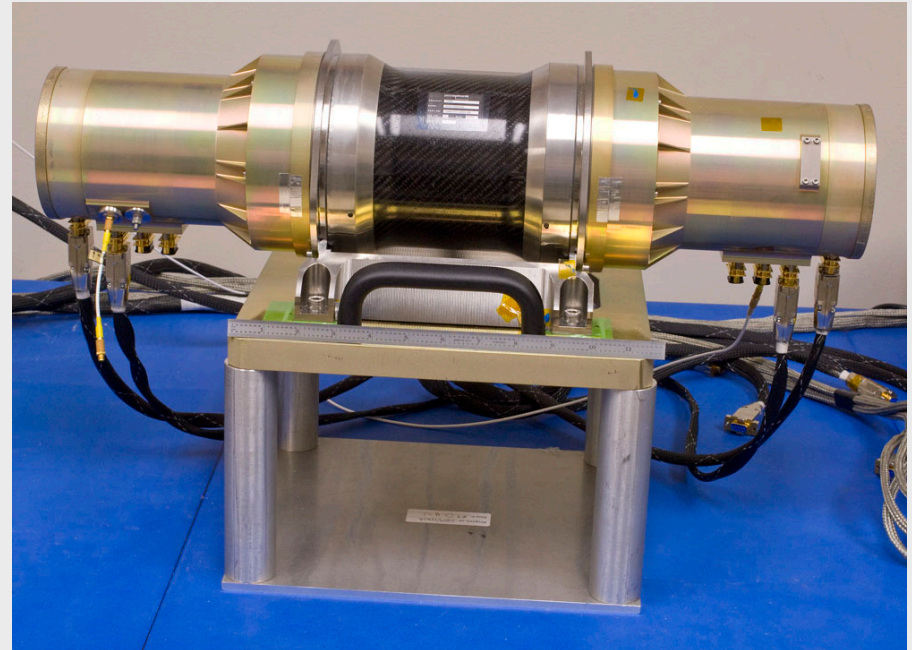








- Sodium iodide (NaI)
- 12.7 cm diameter X 1.27 cm thick
- 8 keV to 1 MeV



- Bismuth germanate (BGO)
- 12.7 cm diameter X 12.7 cm long
- 200 keV to 40 MeV

Very wide field-of-view / non-imaging / high background.

Well suited for transients such as GRBs, solar flares, SGRs, and TGFs.

Persistent or long-timescale sources can be detected by timing analysis (pulsars) or detecting Earth occultation steps.

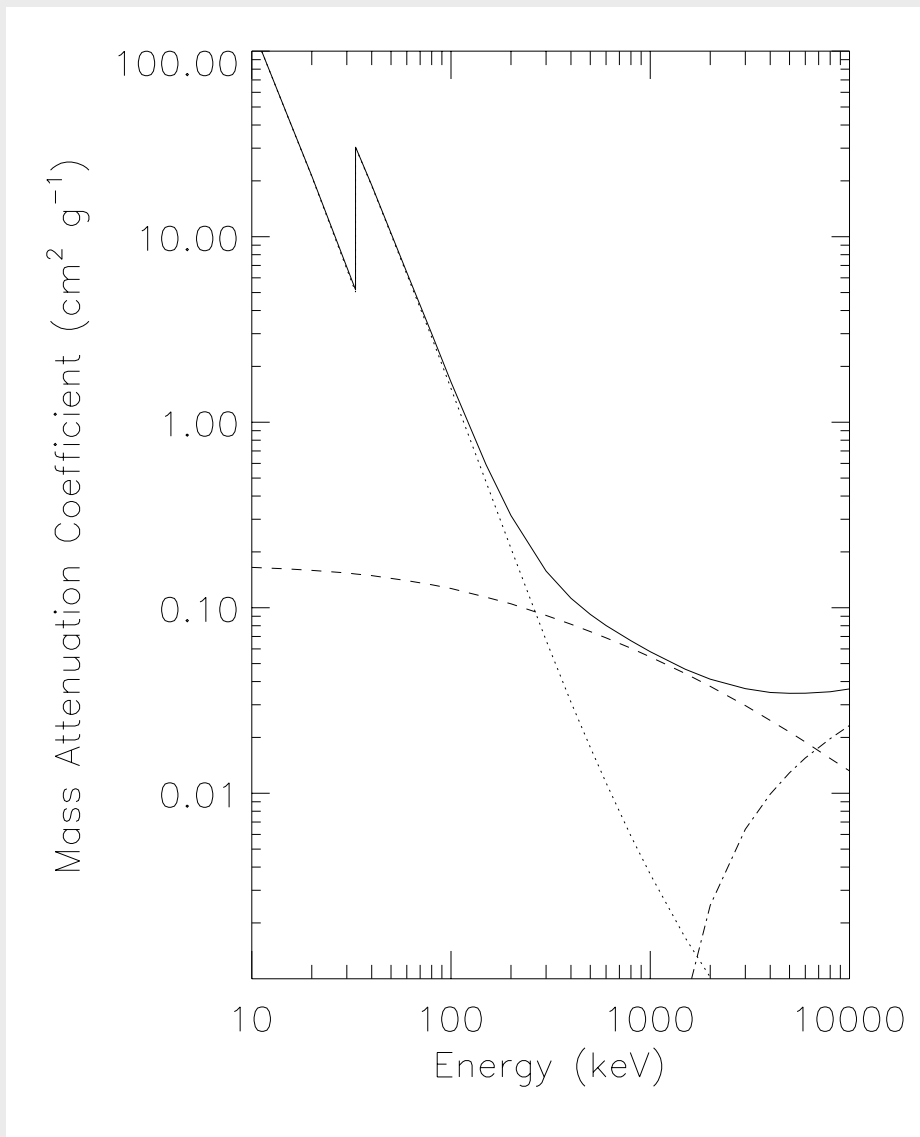


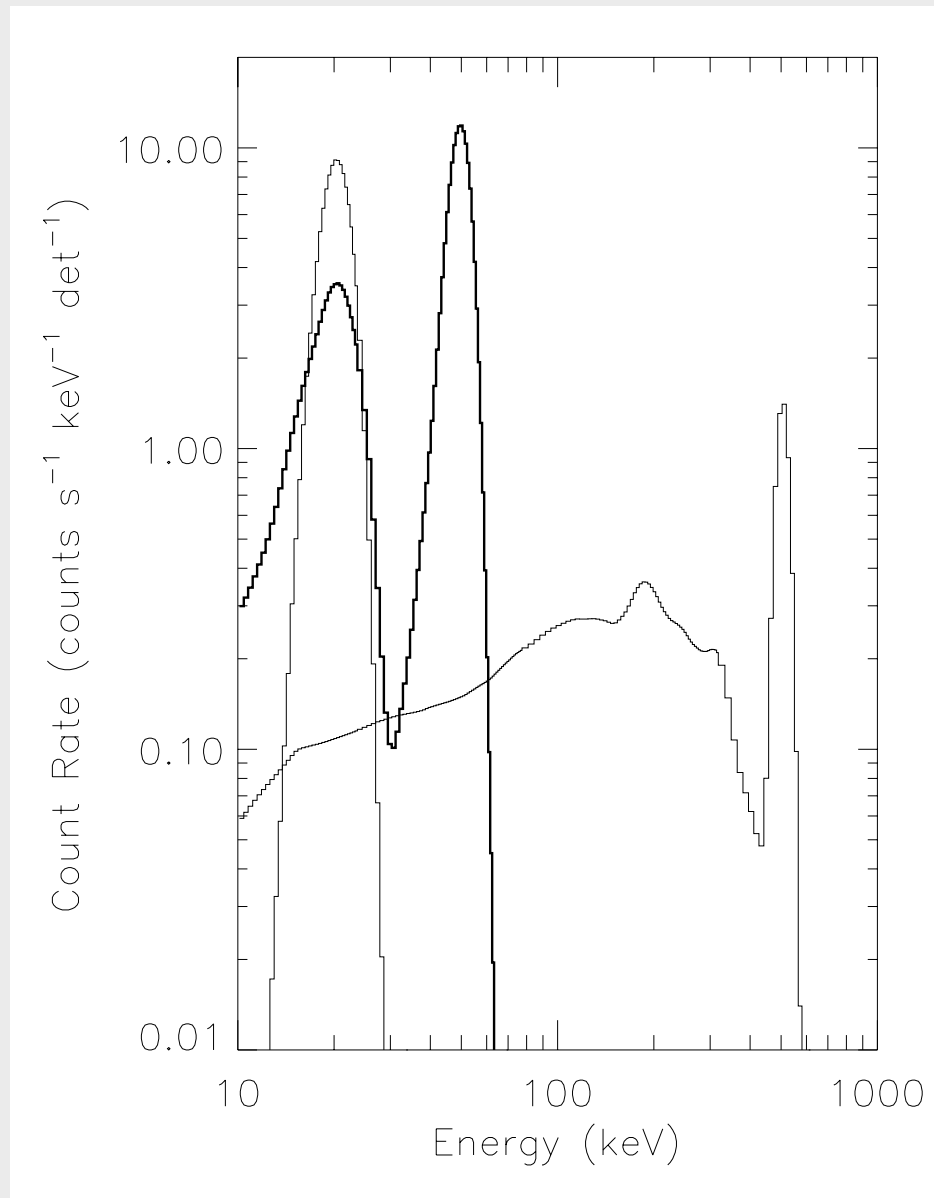
71 AGN are being monitored. There are 12 clear detections:

- Seyfert 2 galaxies:
 - NGC 1275
 - NGC 2110
 - NGC 4388
 - Cen A
 - NGC 5252
 - Circinus Galaxy
 - NGC 5506
- Seyfert 1 galaxies:
 - NGC 4151
 - IC 4329A)
- radio galaxy: IGR J21247+5058)
- quasars:
 - 3C 273
 - 3C 454.3

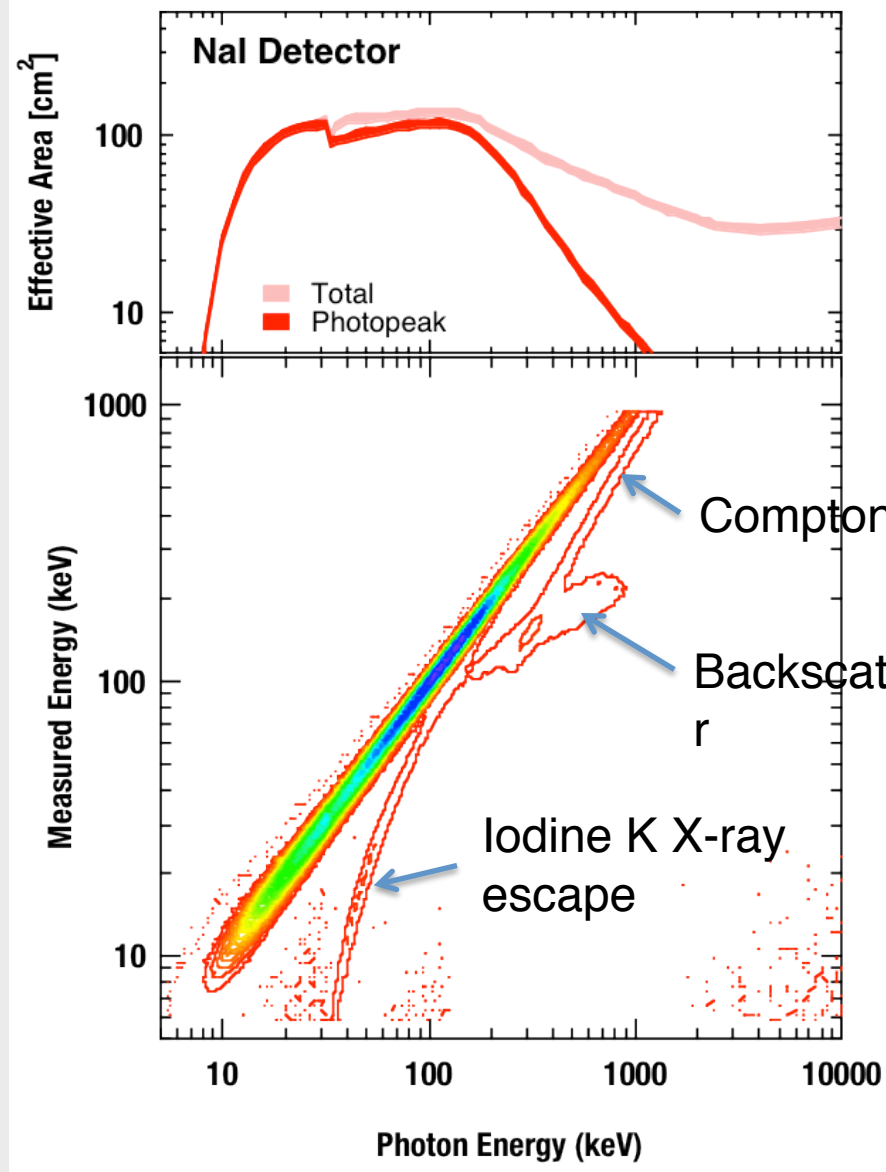


Mass attenuation coefficient μ :
 $T/T_0 = \exp(-\mu x)$,
where
 x is the column density
of NaI traversed
in g cm^{-2} .

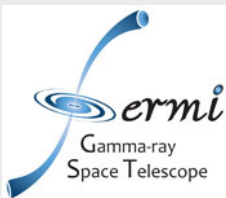
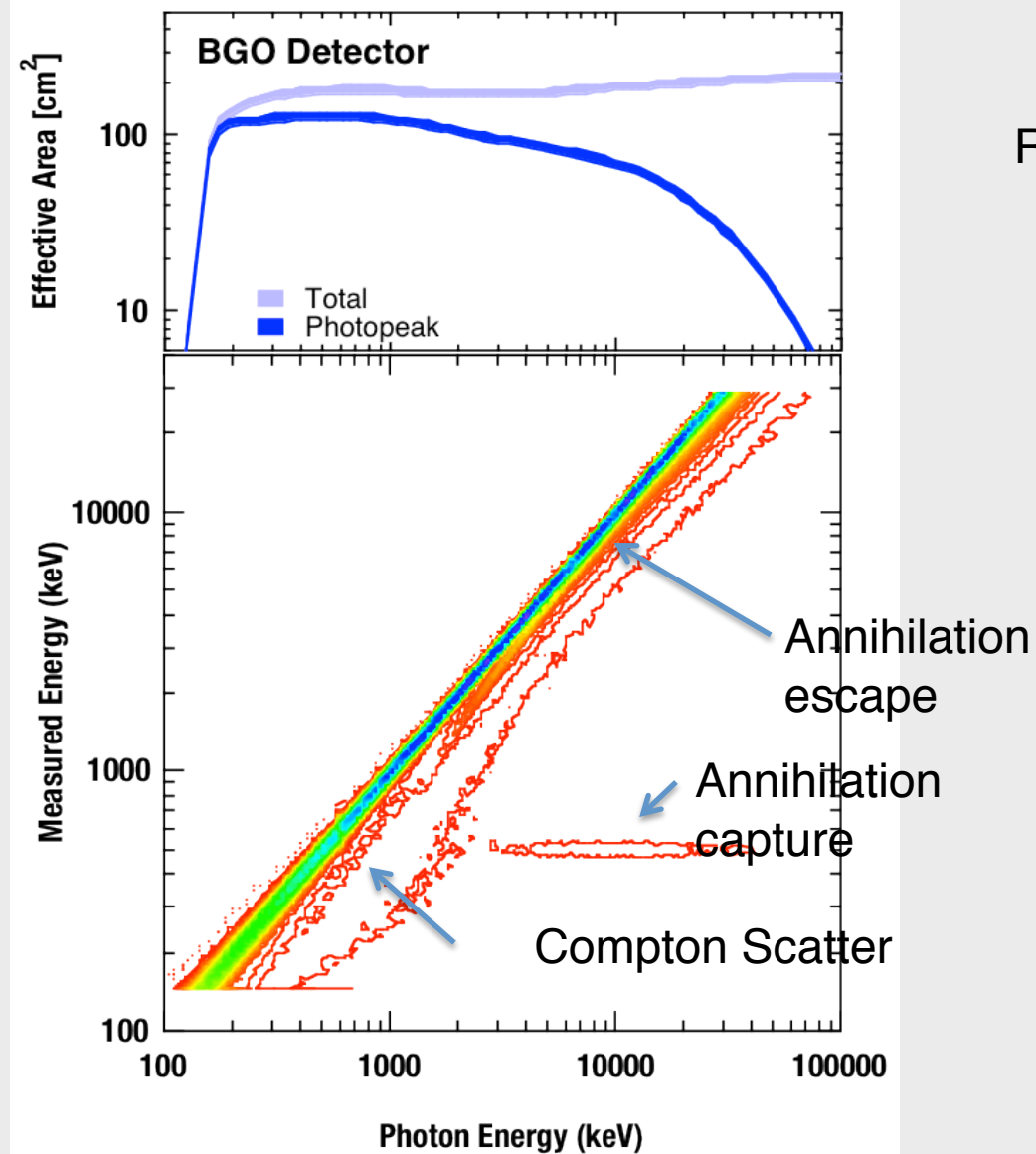




R. M. Kippen



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Count spectrum (binned)

=

DRM X photon model

SO:

Photons = DRM^{-1} X Data ?

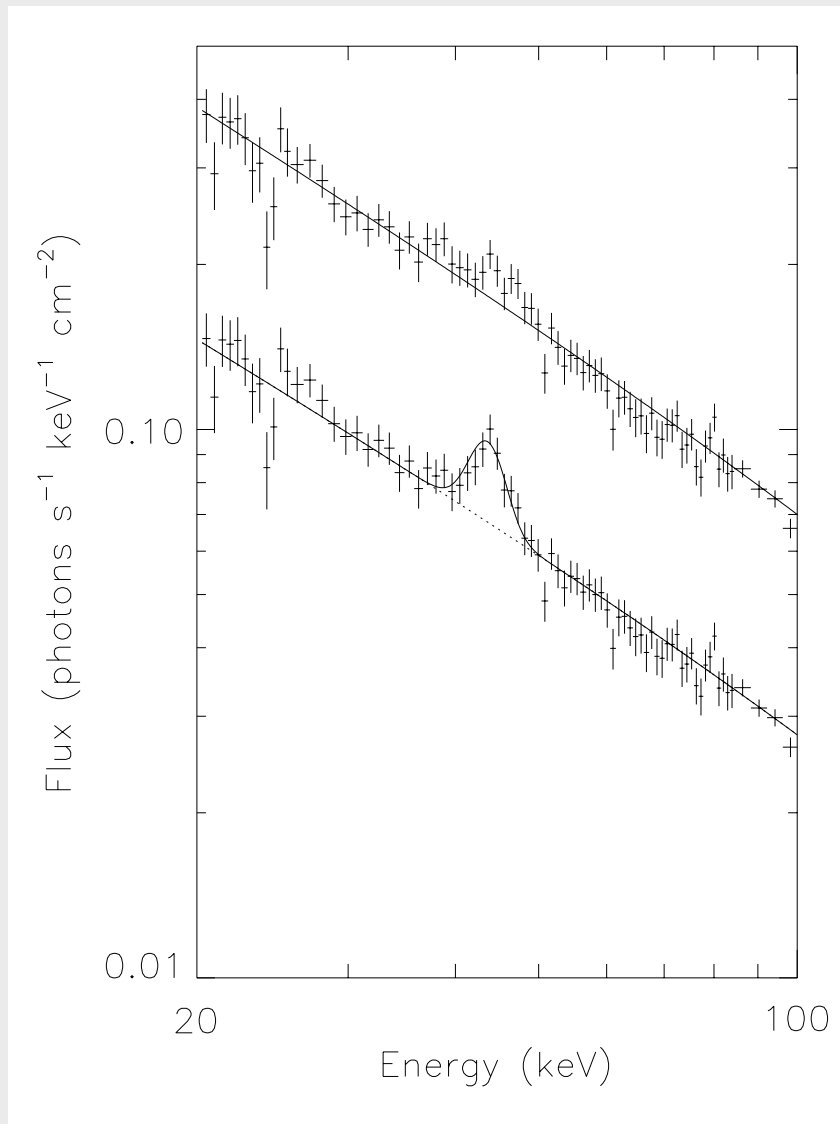


Forward-folding fitting

- 1) Assume a parameterized photon model.
- 2) Select a fitting statistic – likelihood or C-Stat (or χ^2)
- 3) Calculate the count model using the DRM / IRF,
- 4) Calculate the fitting statistics,
- 5) Change the photon model parameters to improve the fitting statistic,
- 6) Repeat steps 3 to 5 to optimize the model.
- 7) The answer is based upon the model that you assumed – the process cannot automatically find the true model.



What not do do!

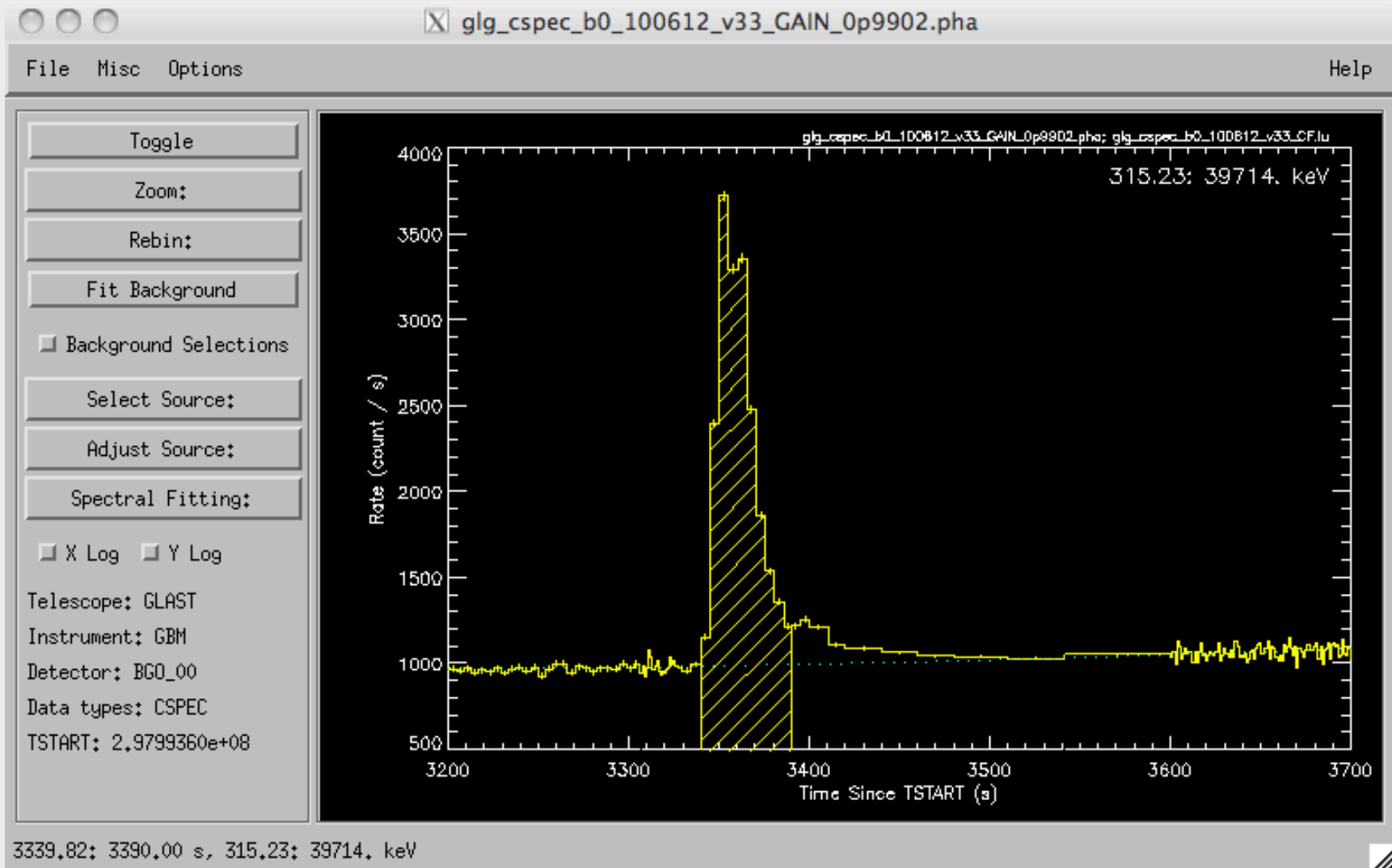


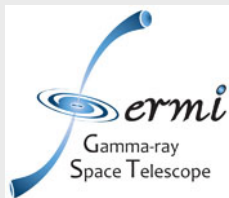
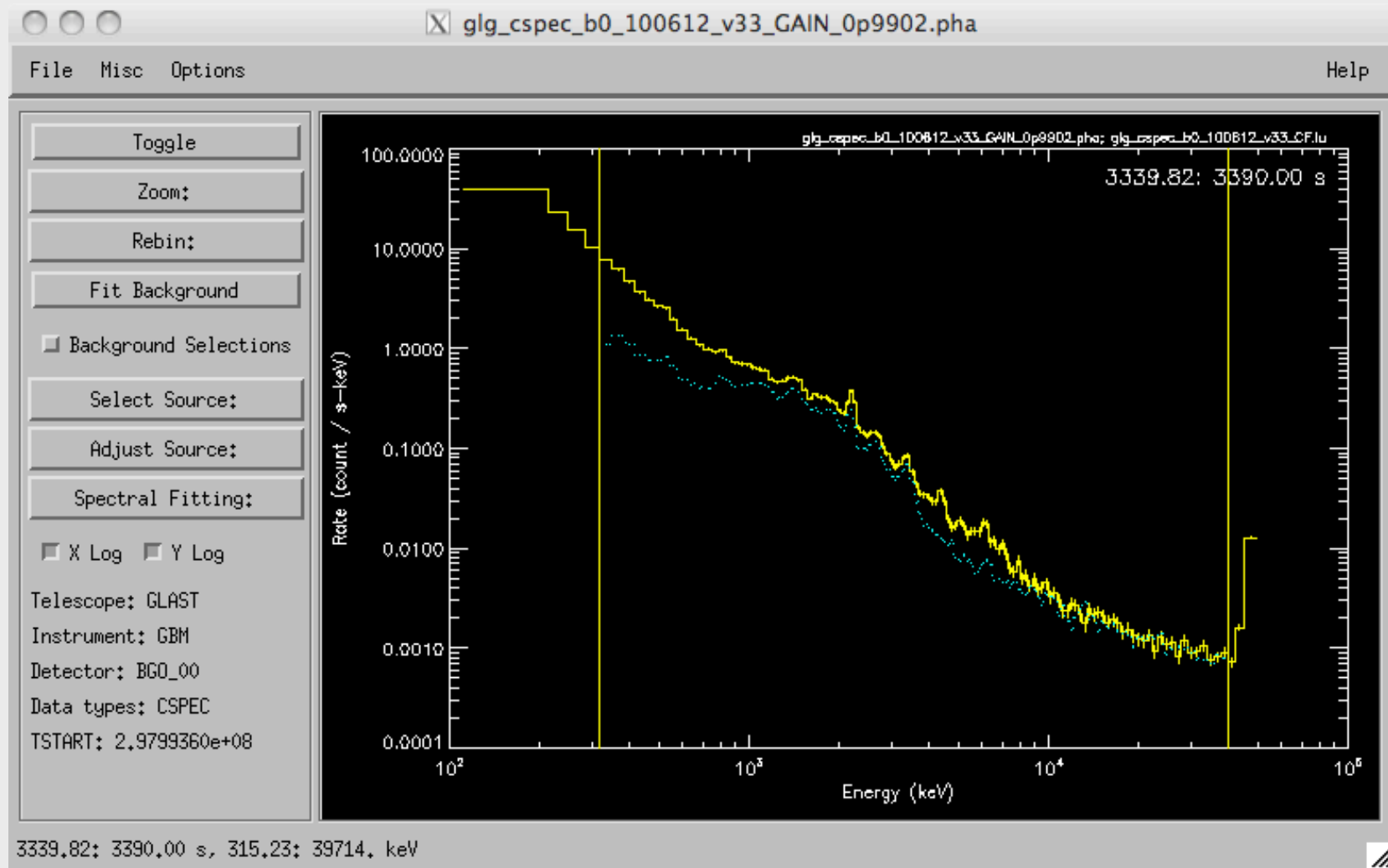
GBM Datatypes

GBM datatypes: designed to meet a stringent telemetry requirement. Currently 2% of the mission volume.

- Triggered:
 - CSPEC: 128 channels, binned at 1.024 s
 - CTIME: 8 channels, binned at 0.064 s
 - TTE: 128 channels, individual counts at 2 μ s
- Continuous / daily:
 - CSPEC: 128 channels, binned at 4.096 s
 - CTIME: 8 channels, binned at 0.256 s
 - **NEW**: TTE: 128 channels, individual counts at 2 μ s







PhotonModel

Select one or more photon model terms

- 1 Power Law
- 2 Power Law B
- 3 Broken Power Law
- 4 Power Law w. 2 Breaks
- 5 Smoothly Broken Power Law
- 6 Band's GRB, Epeak
- 7 Band's GRB, Epeak, B
- 8 Old Band's GRB
- 9 Comptonized, Epeak
- 10 Comptonized, Epeak, B
- 11 Old Comptonized
- 12 Brainerd's Scattered Pow Law
- 13 Log Normal
- 14 Gaussian (log10 E)
- 15 Gauss (log10 E), linear FWHM

Photon Model Parameters:

Keep current Set parameters

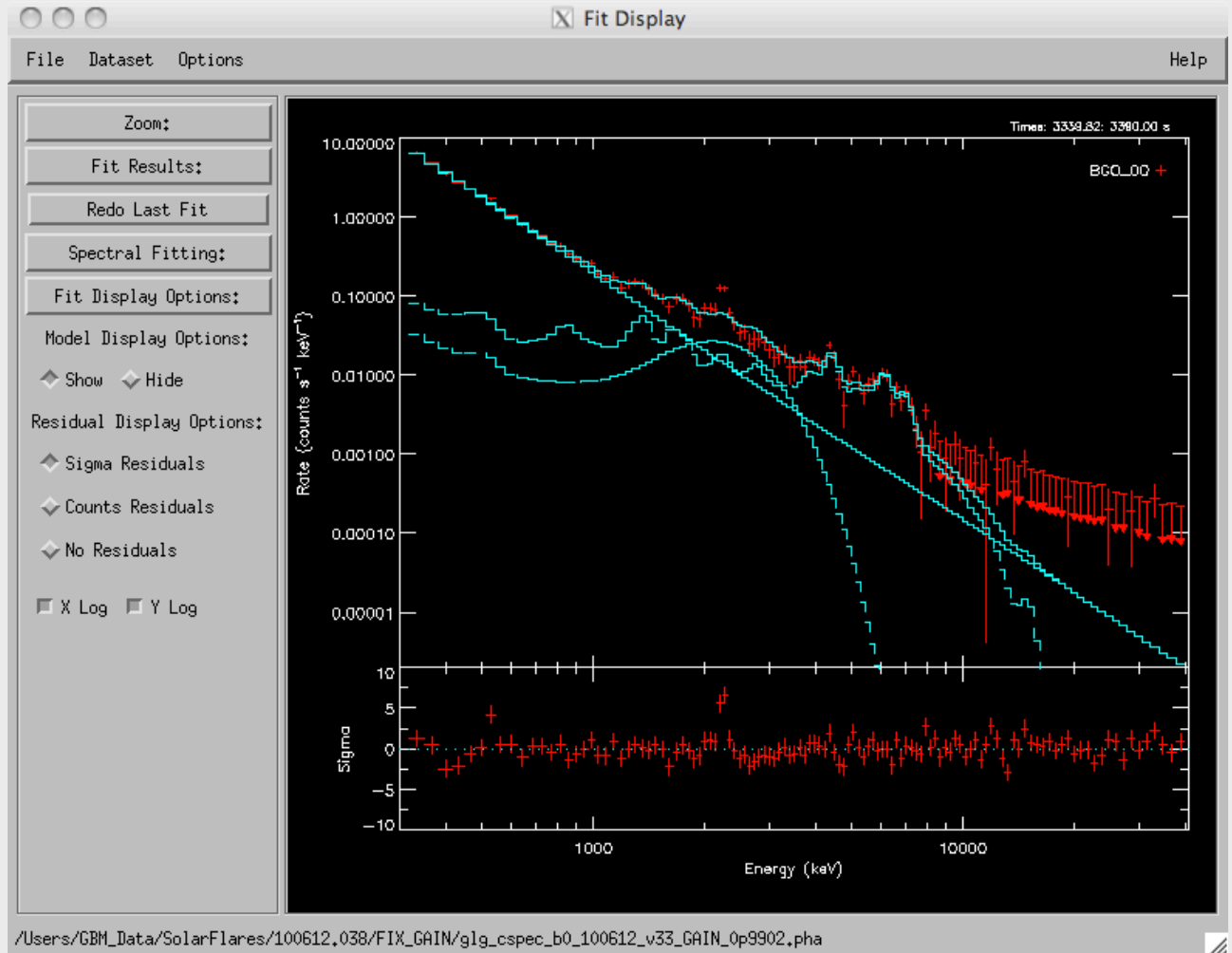
Fitting Statistic:

Chi² Likelihood C-Stat

Undetermined Values in Batch Fit:

Leave Free Automatically Fix

Accept Omit Restore Cancel



Zoom:

Fit Results:

Redo Last Fit

Spectral Fitting:

Fit Display Options:

Model Display Options:

Show Hide

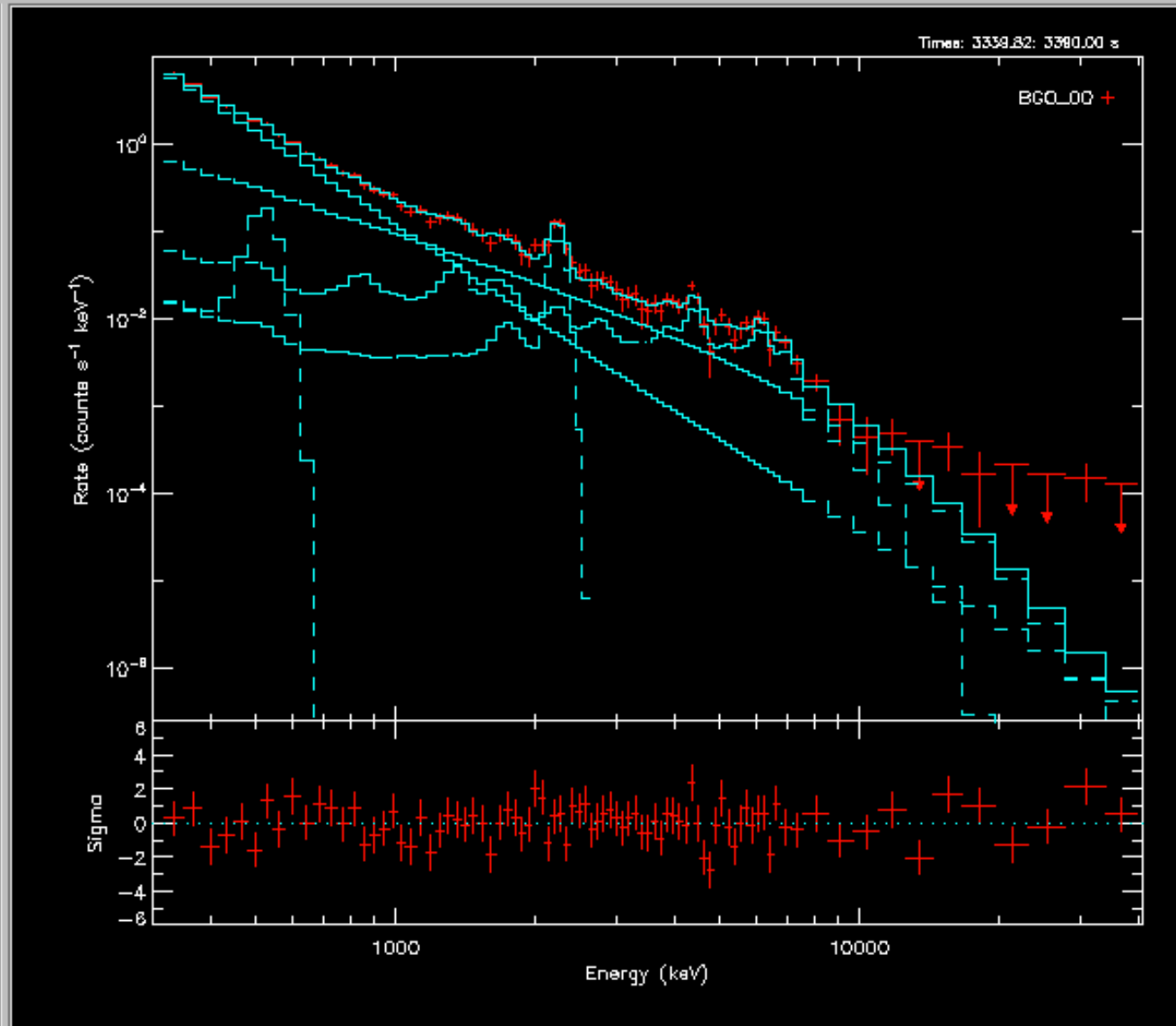
Residual Display Options:

Sigma Residuals

Counts Residuals

No Residuals

X Log Y Log



/Users/GBM_Data/SolarFlares/100612.038/FIX_GAIN/glg_cspect_b0_100612_v33_GAIN_0p9902.pha

Fit Display

File Dataset Options

Help

Zoom:

Fit Results:

Redo Last Fit

Spectral Fitting:

Fit Display Options:

Model Display Options:

Show Hide

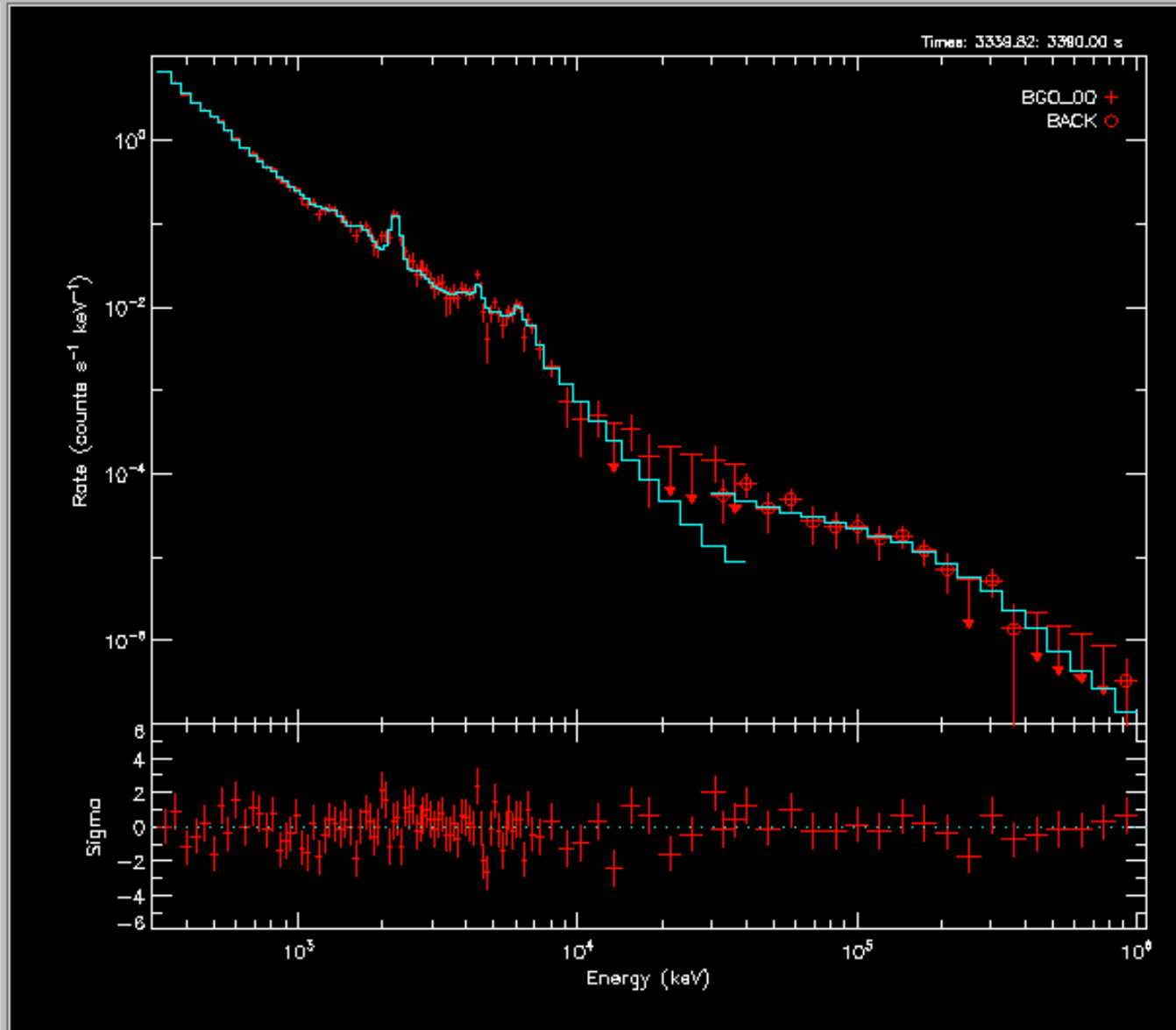
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Zoom:

Fit Results:

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Fit Display Options:

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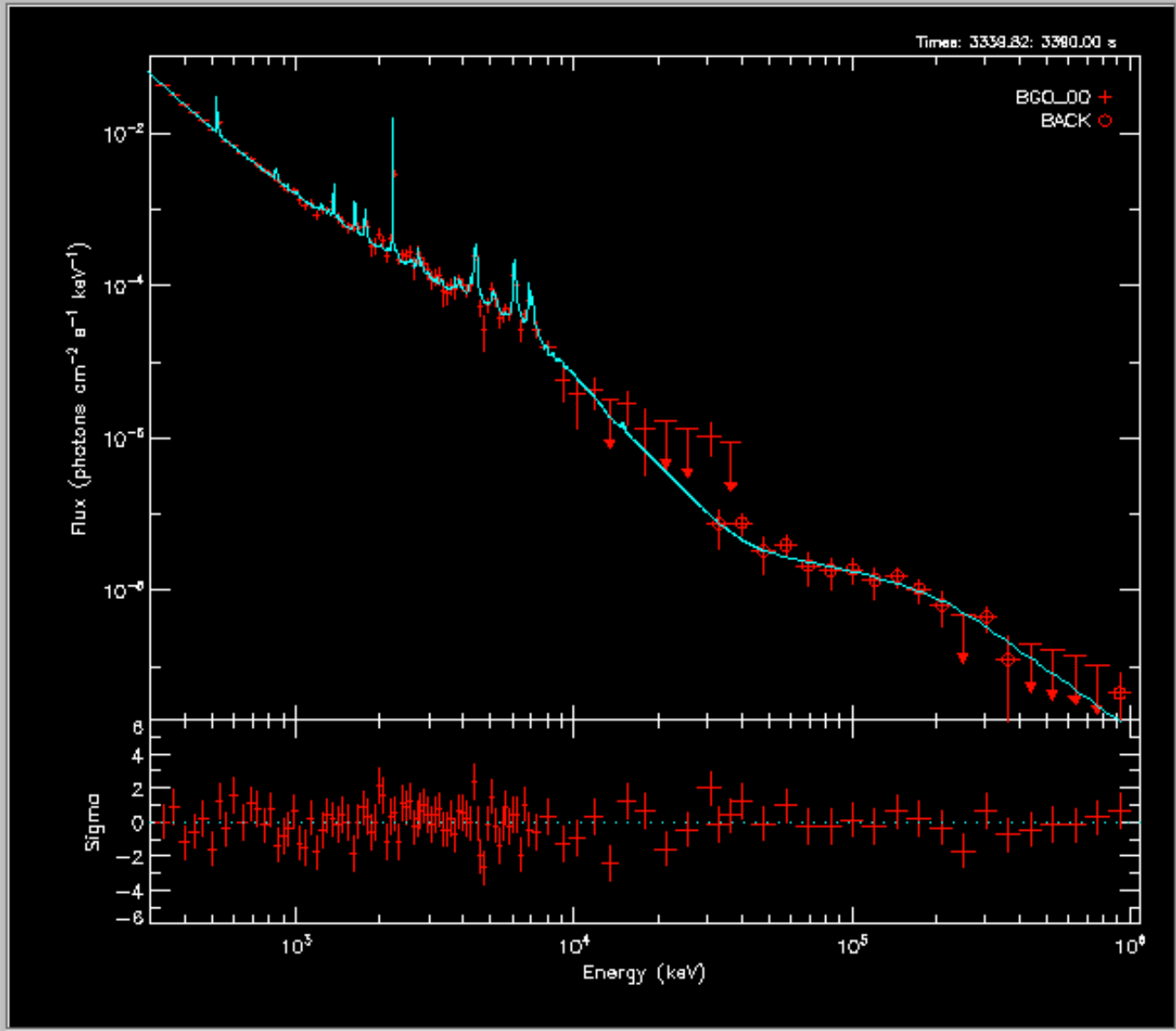
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