EGRET discovery of Gamma-ray blazars

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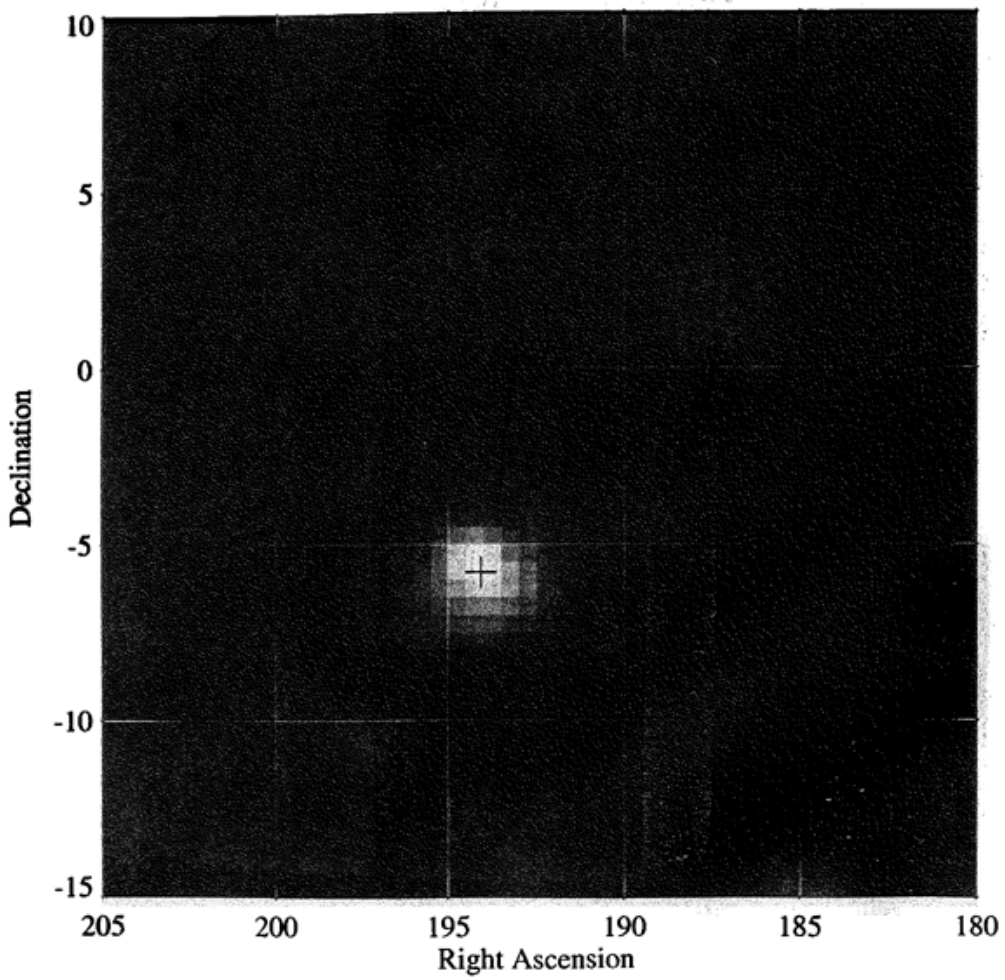
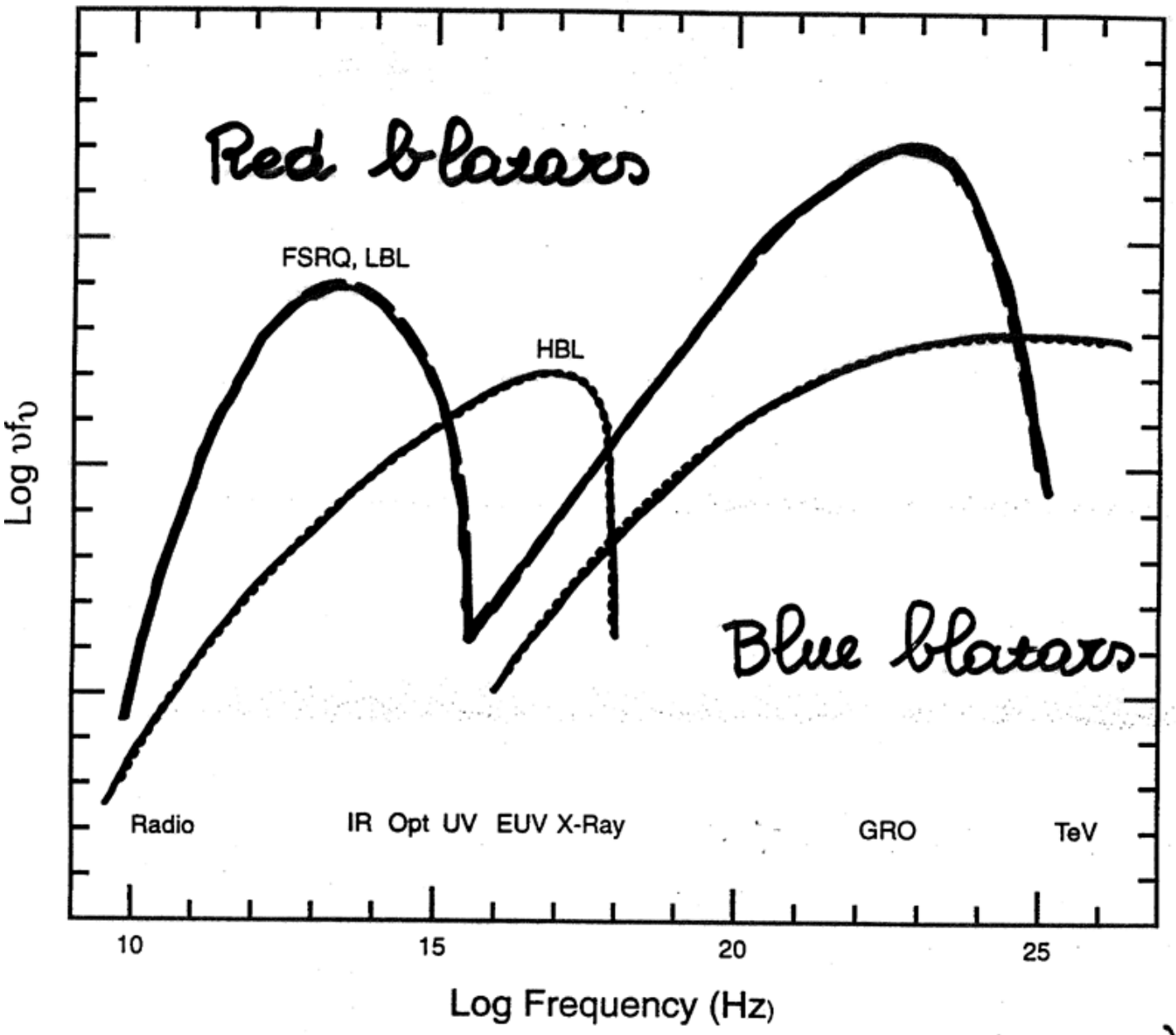


3C 273

3C 279

RMS



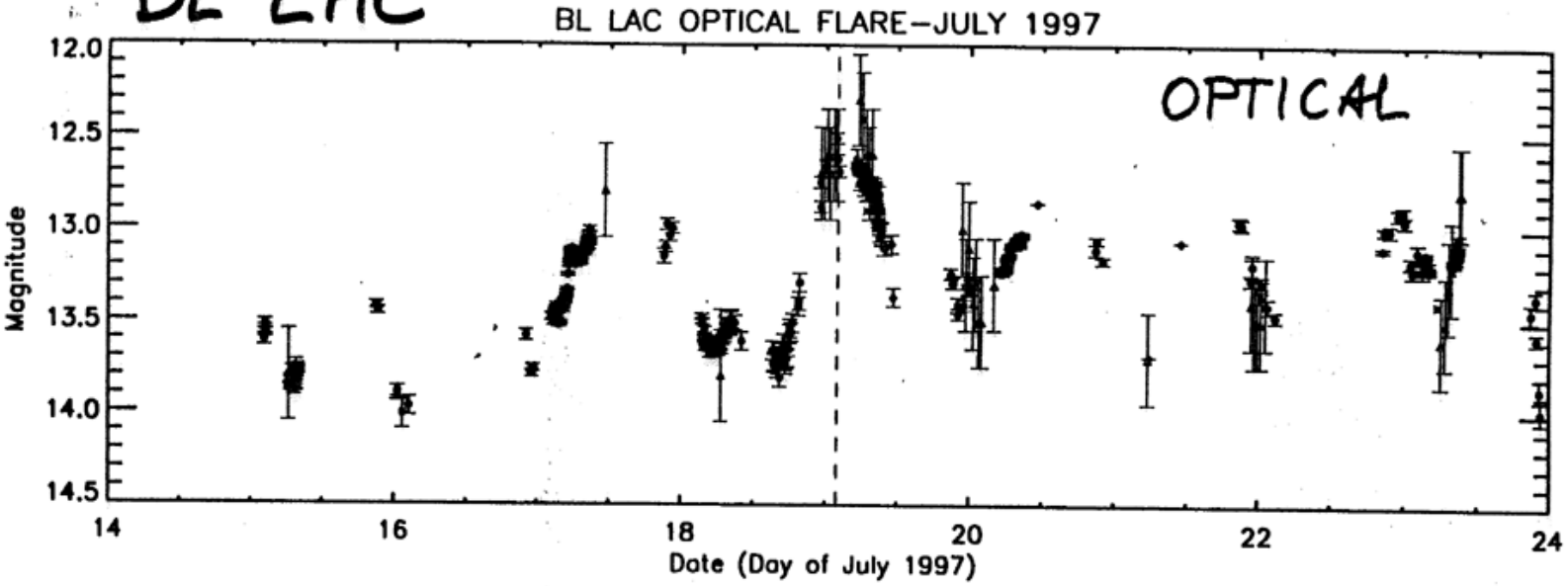
- synchrotron + ? inverse Compton (SSC, EC)

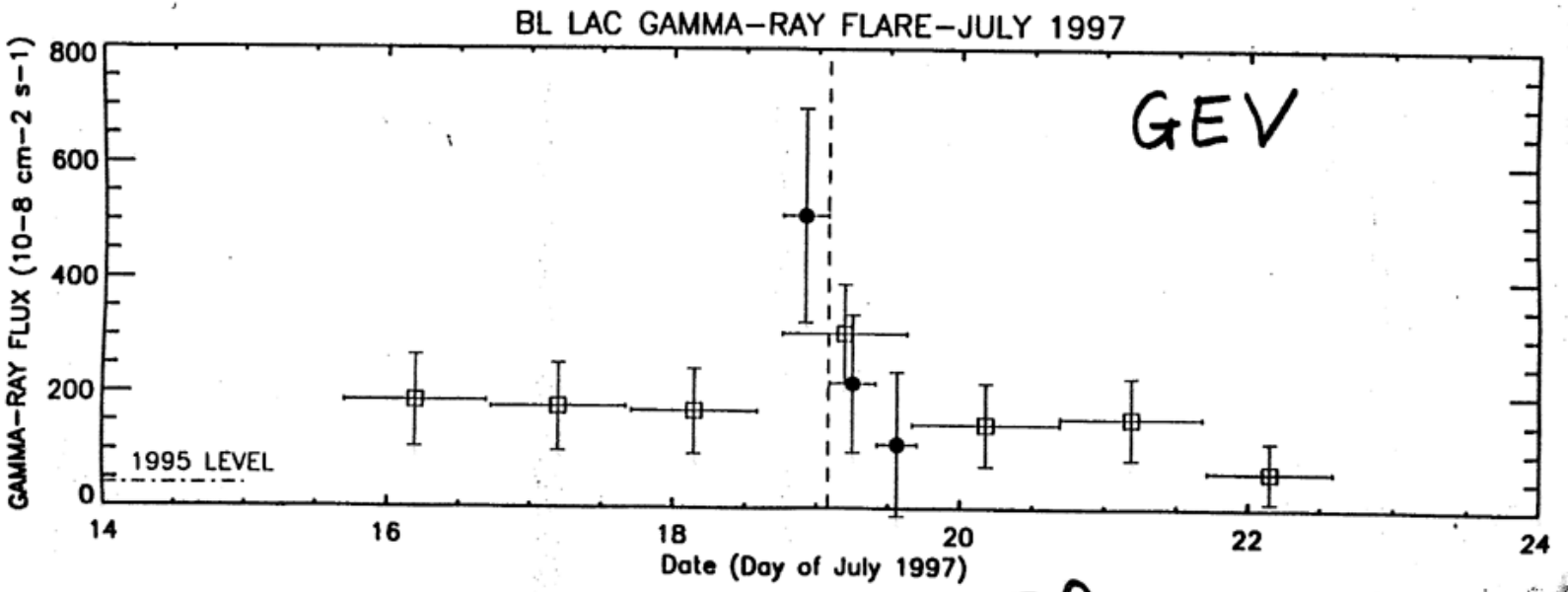
PIC

RMS

Date

BL LAC

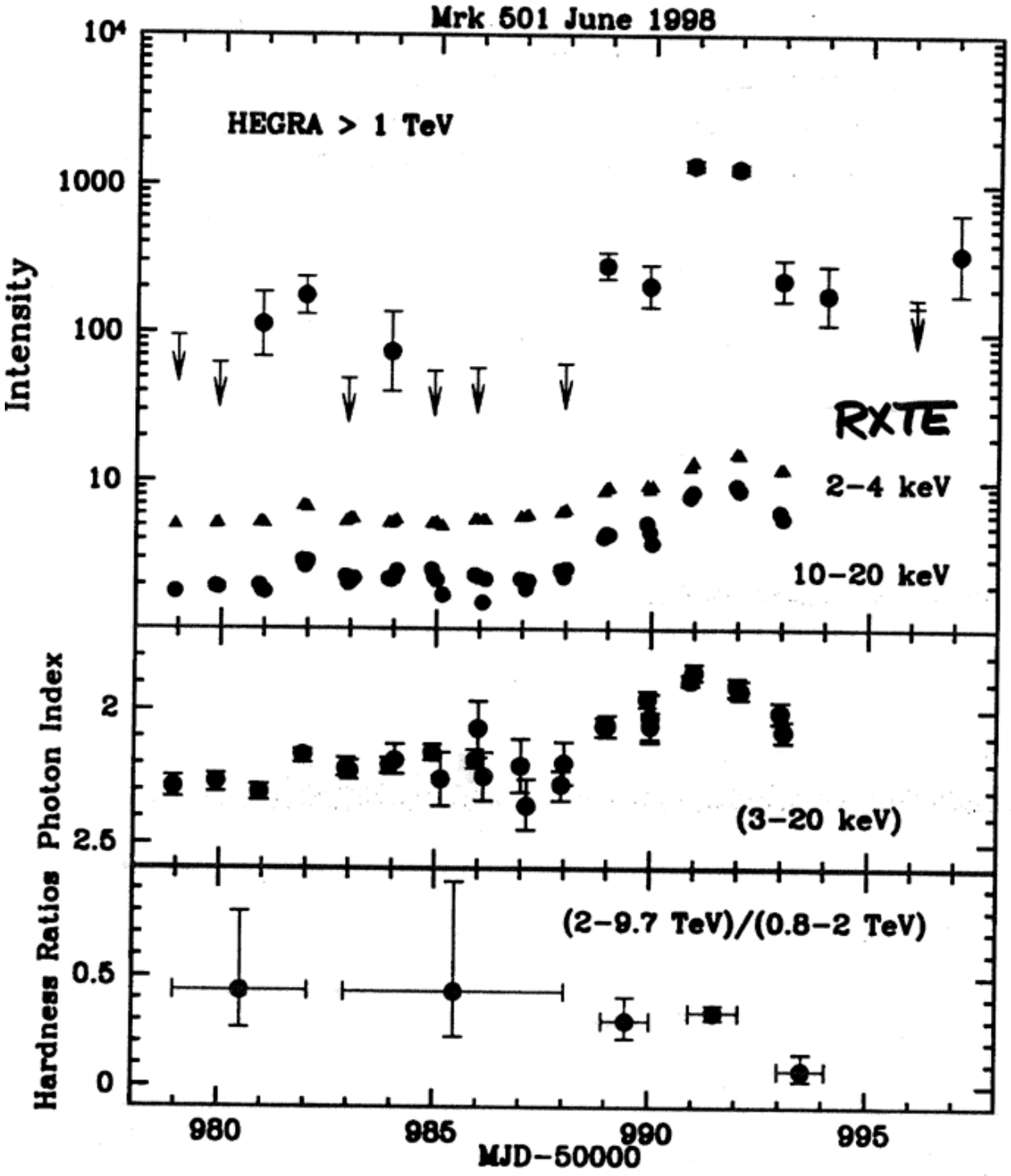




KTS

Bloom et al. 1998

RHS



SamBruna + 2000

Synchrotron peak frequency

$$\nu_s \propto \gamma^2 B$$

IC peak frequency

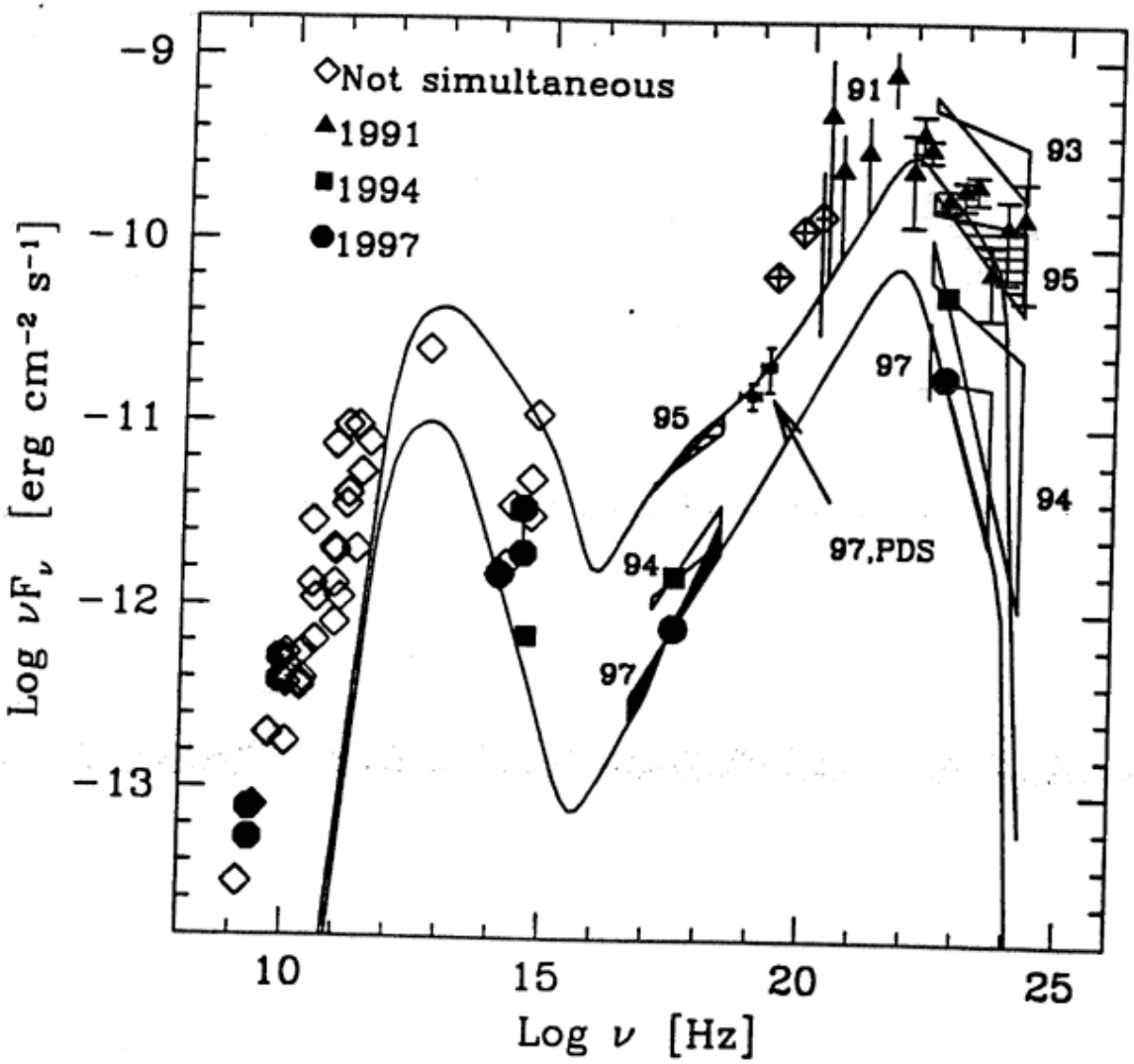
$$\nu_{IC} \propto \begin{cases} \gamma^4 B & \text{SSC} \\ \gamma^2 \nu_{Ext} & \text{EC} \end{cases}$$

1) $$\frac{\nu_{IC}}{\nu_s} \propto \begin{cases} \gamma^2 & \text{SSC} \\ \frac{\nu_{Ext}}{B} & \text{EC} \end{cases}$$

2) $$\frac{L_{IC}}{L_s} \propto \frac{U_{en,photons}}{U_B}$$

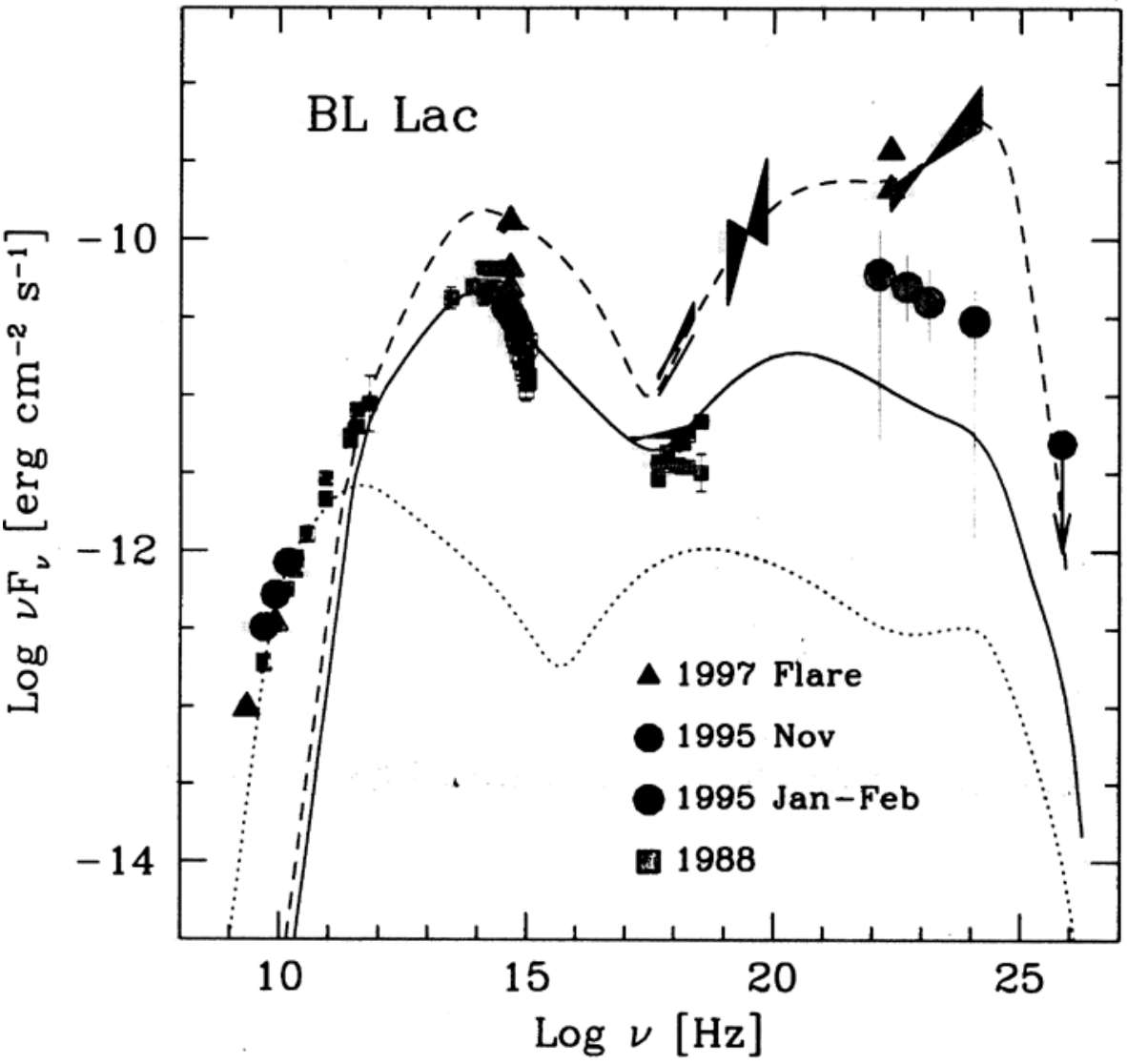
RMS

PKS 0528+134



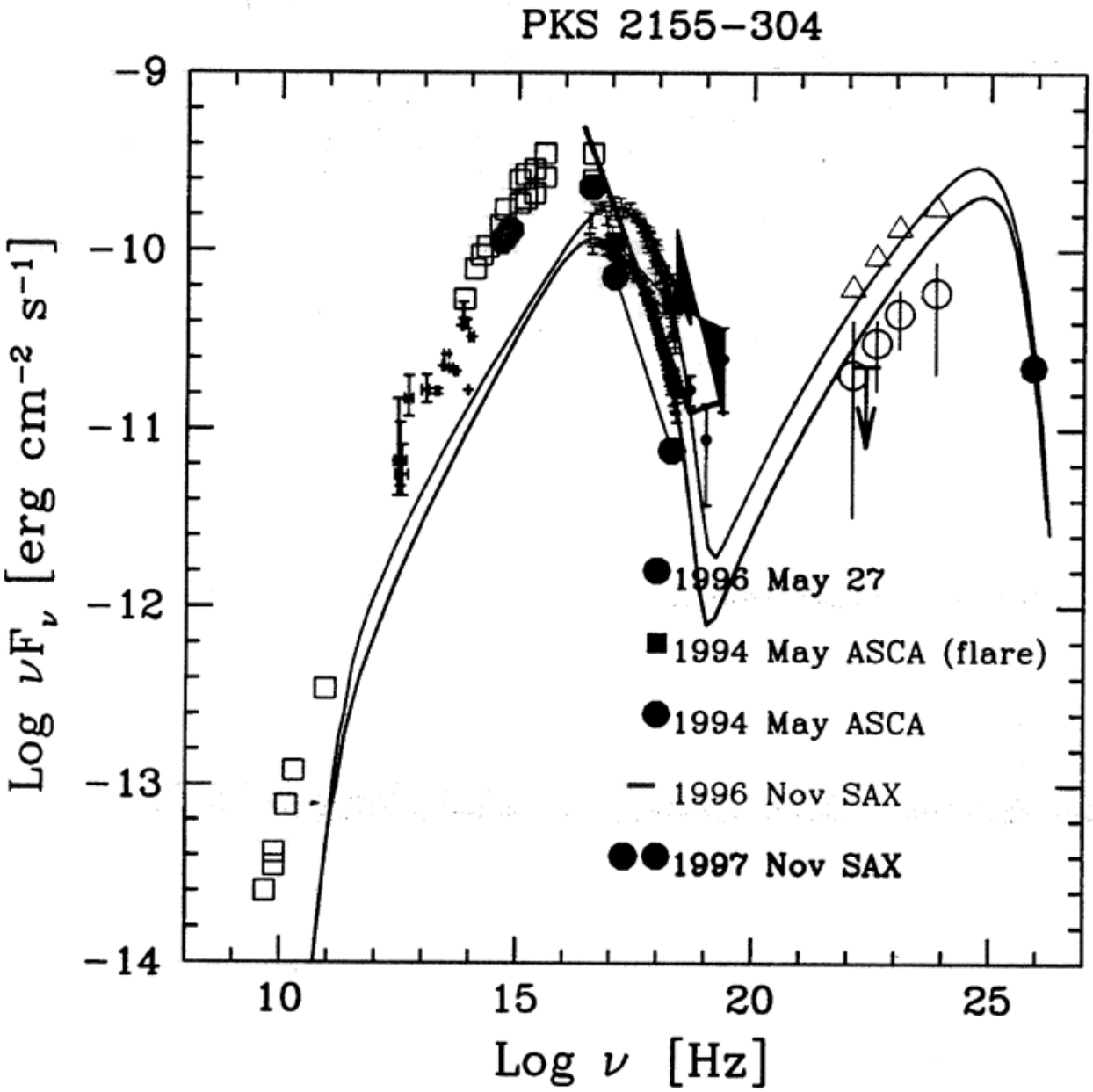
Ghisellini + 98

RHS



Sambruna et al. 98

RHS

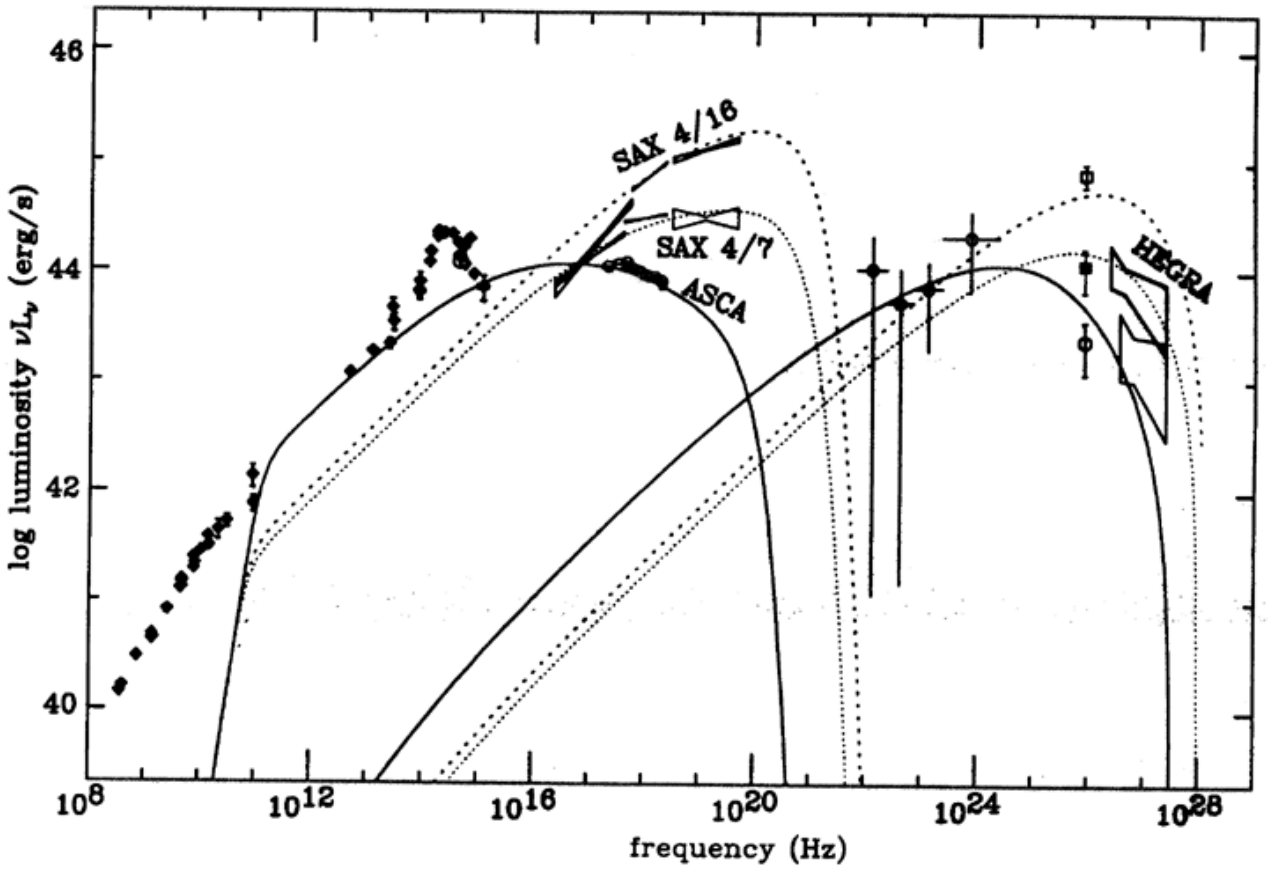


Bedtone + 1999

● ● April 1997
● March 1996
● Literature

RHS

Mrk 501



Kataoka + 1999

RMS

Essential λ bands

	ν_s	ν_{IC}
Red bl.	IR OPT	keV - GeV Integral, Astro GLAST
Blue bl.	UV X	MeV - TeV GLAST HESS, VERITAS, MAGIC, CELESTE, STACEE,

Cross - Calibration important !

RMS

GLAST observations of blazars

Red bl

curved γ spectra

large γ vblty @ higher E

Blue bl

flatter γ spectra

lower γ vblty.

γ Synchro + IC

How will we use GLAST? RMS

Conservative approach:

"famous" targets (3C 279, 0528, ...

intensive monitoring

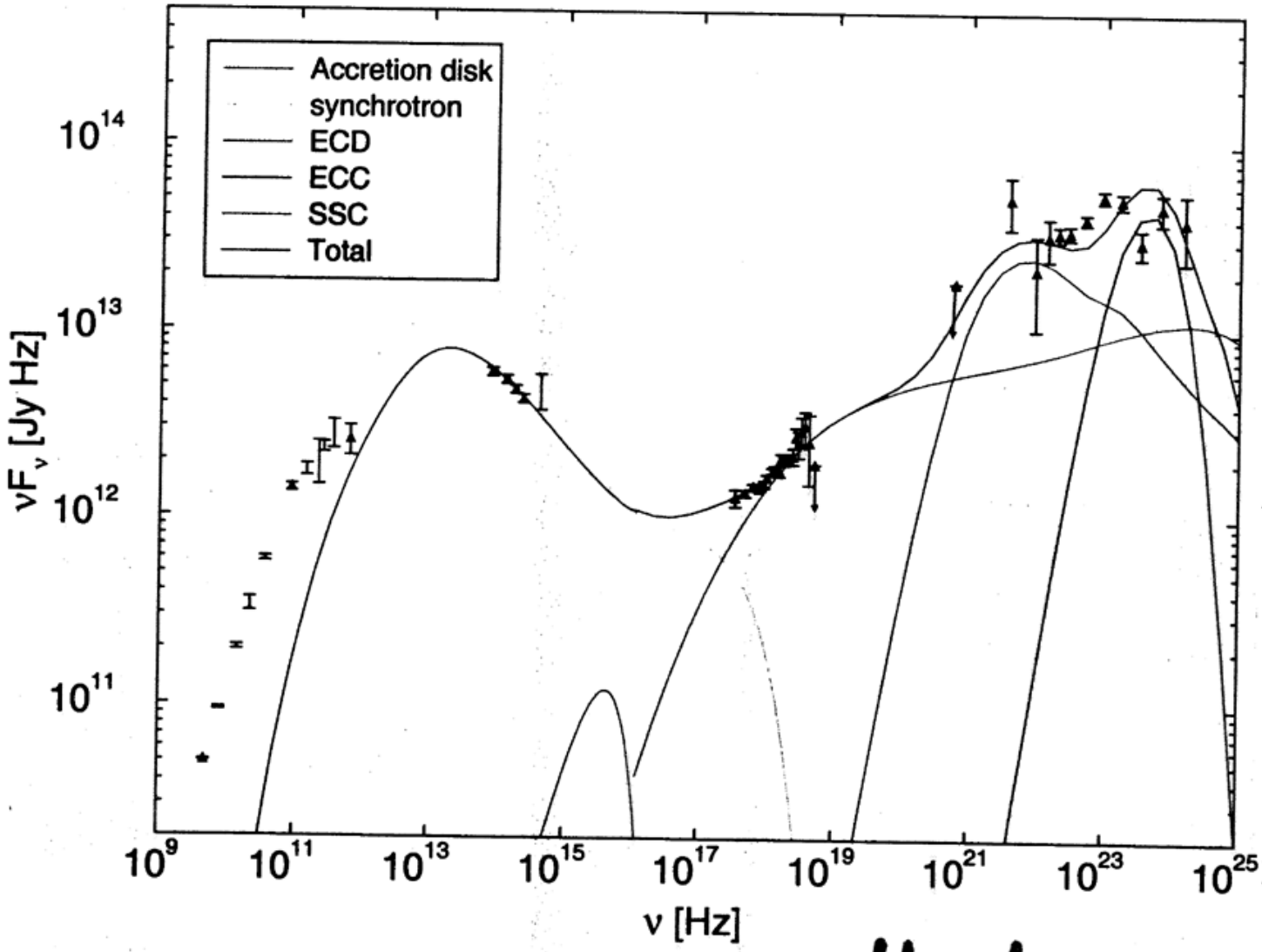
high / low states

- Origin of γ → detailed γ + MW spectra
- structure of jet → γ flare shape MW correlations
- particle dynamics → flux / spectral variability

...

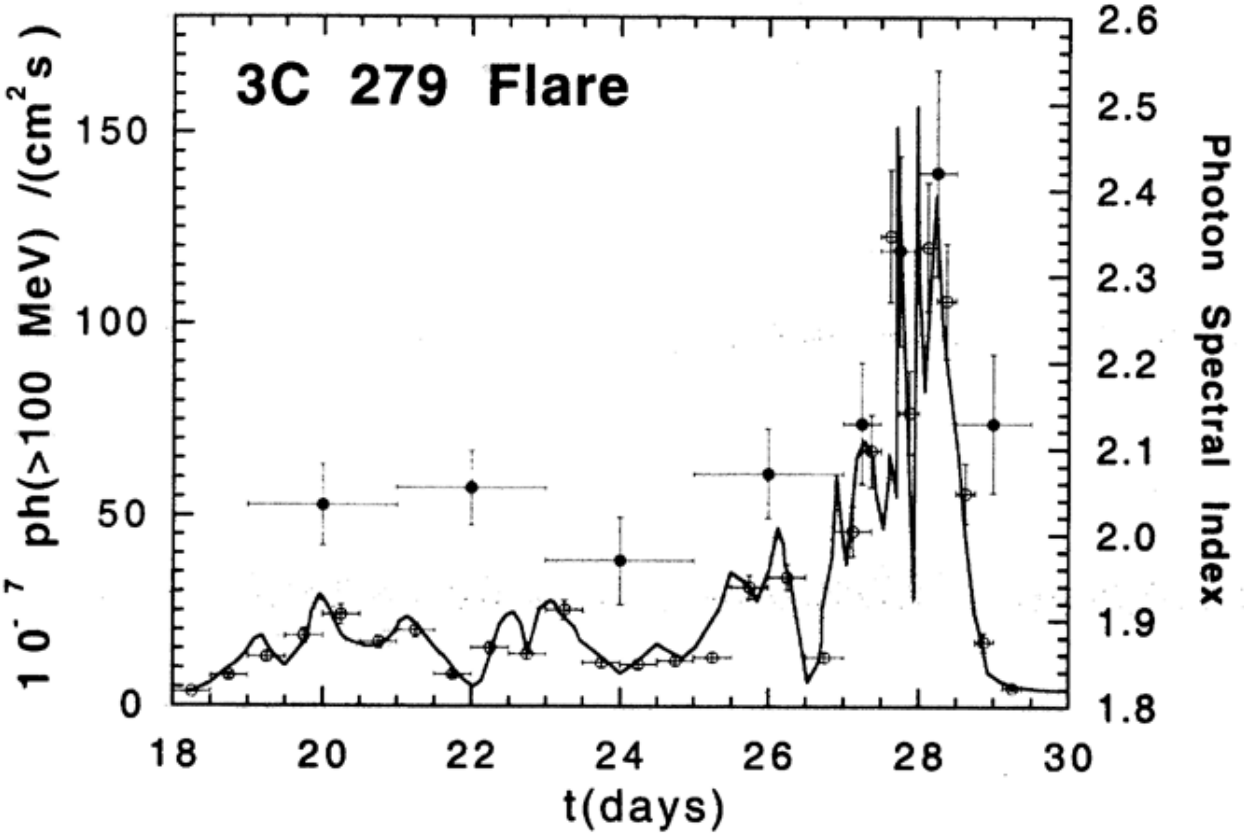
RHS

3C 279 1991 June, flare state (period P1)



Hartman + 2000

RHS


3C 279 Flare
10^-7 ph(>100 MeV) /(cm^2 s)
Photon Spectral Index
t(days)
0
50
100
150
1.8
1.9
2.0
2.1
2.2
2.3
2.4
2.5
2.6
18
20
22
24
26
28
30

(C. Dermer)

RMS

Cover as many blazars as possible

- new behaviors

- calibrate $\frac{L_{IC}}{L_S}$, $\frac{\nu_{IC}}{\nu_S}$

- duty cycle of γ → are all blazar γ emitters?

Large FOV

⇒ many sources "for free"

GLAST obs triggered by other λ

OPT	RED BL.
X	BLUE BL.

RMS

GLAST as a trigger for ToO ?

Depends:

- How fast info transmitted to ground
- How often
- How quick the "quick look" is
- the type / duration of flares

$t_\gamma \sim 1\,d$

$\sim$ few hrs

(less ?)

RMS

"safest" approach

plan λλλ coverage

<u>at least</u> in critical

bands

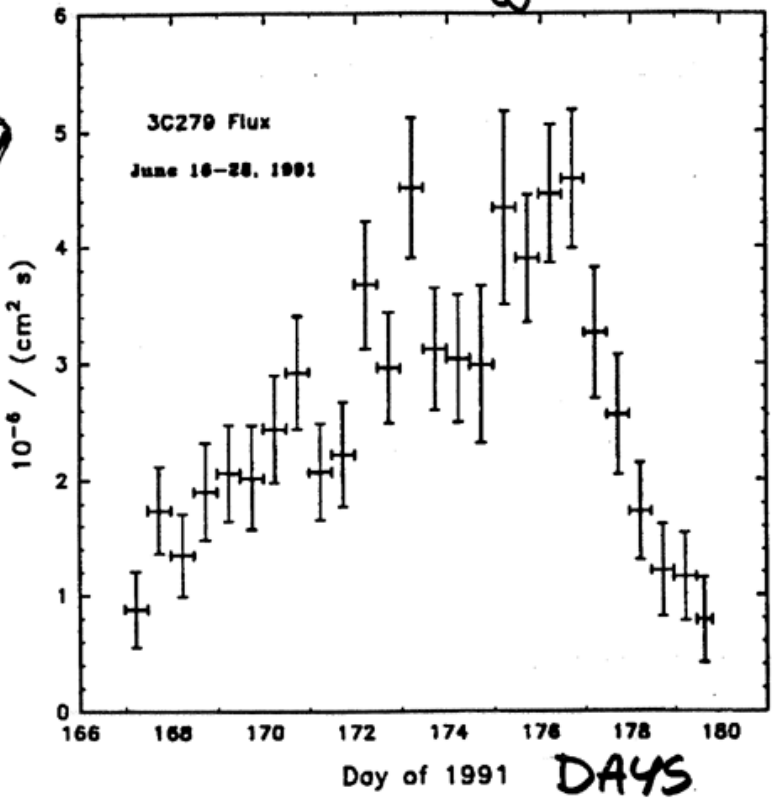
... don't discard

upper limits!

RMS

Knuffen et al. 1993

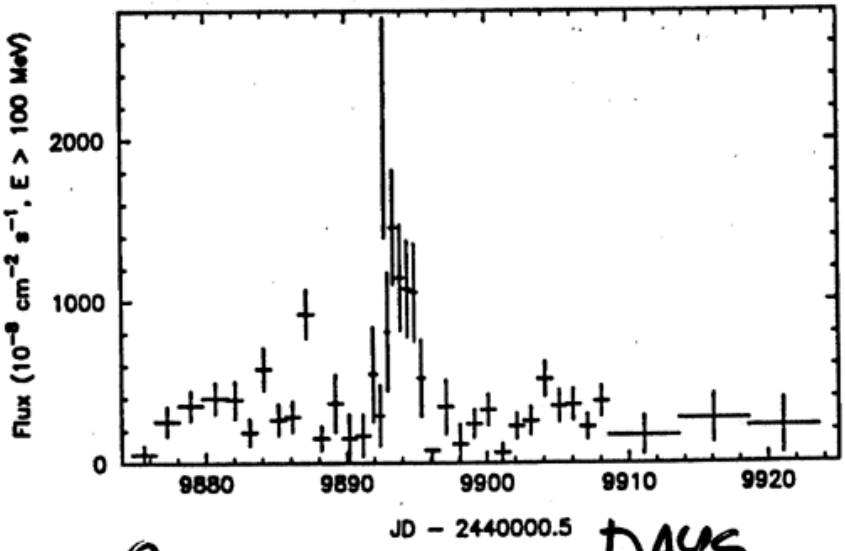
3C 279



$t_y \sim 1$ d

PKS 1622-297

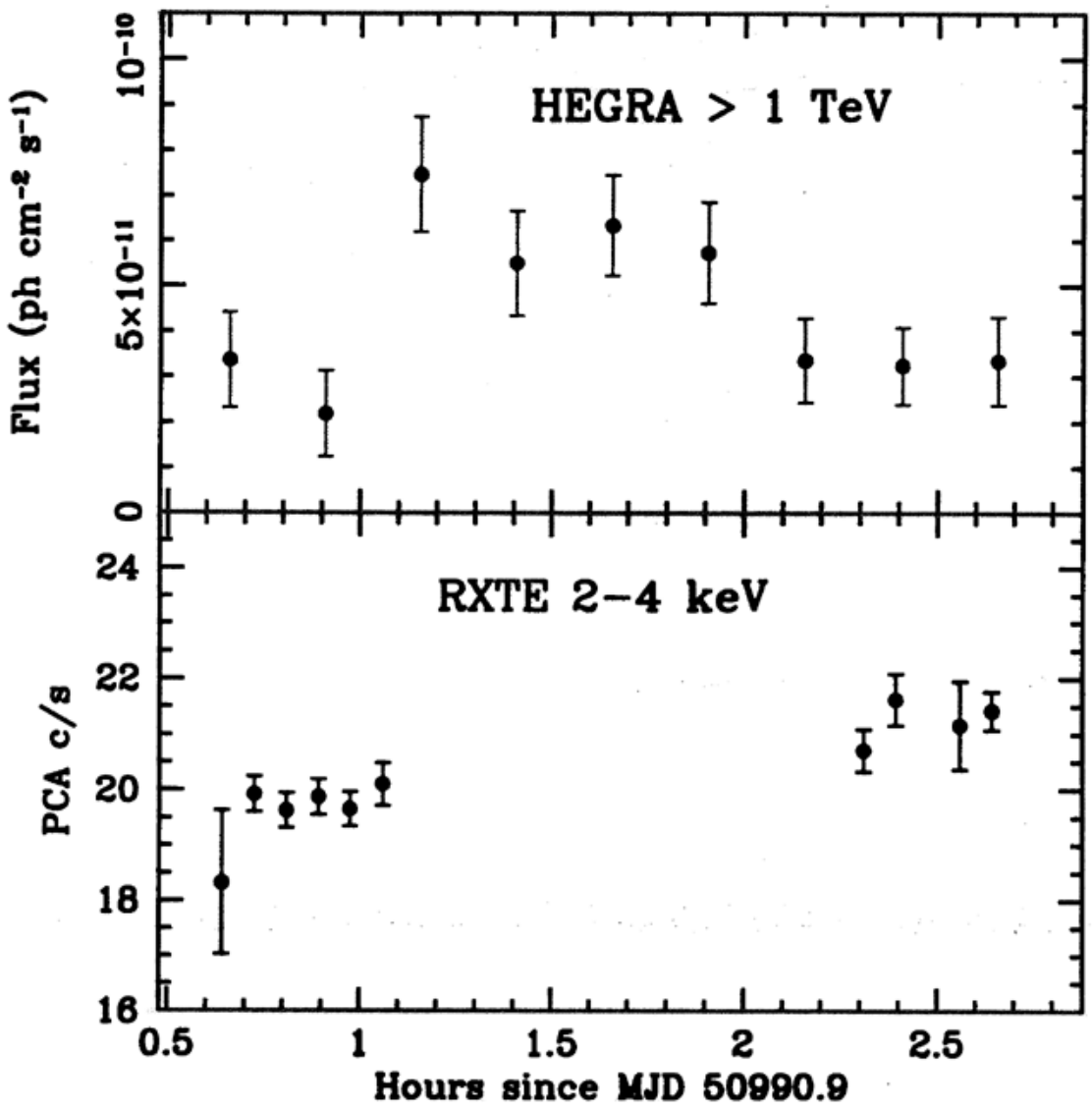
Mattox et al. 1997



$t_d \sim 3.8$ Rgs

Sub-hour TeV variability in M501

RMS



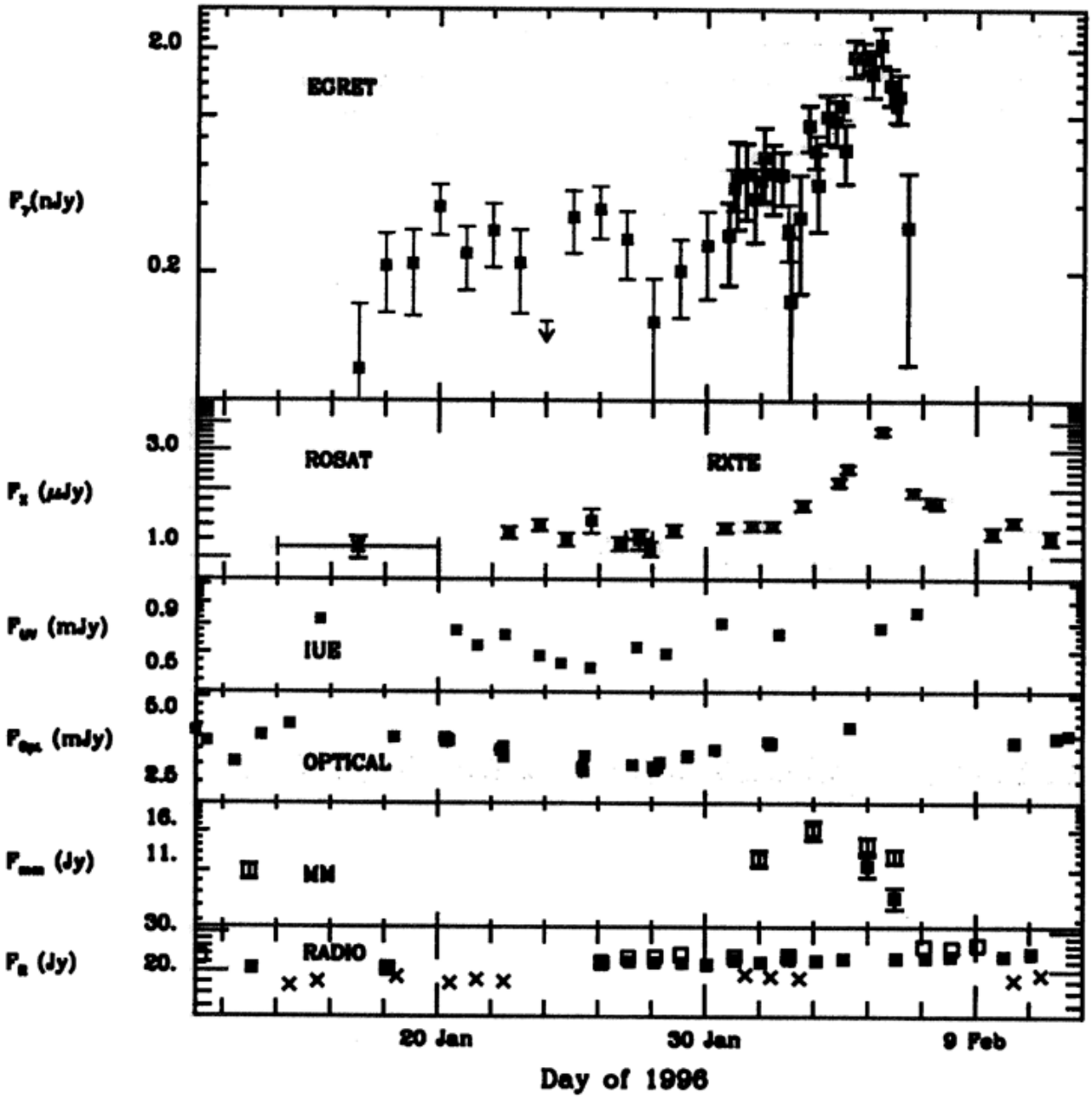
$t_d \sim 20$ min

$R \lesssim 10^{16}$ cm

R. Sambruna +

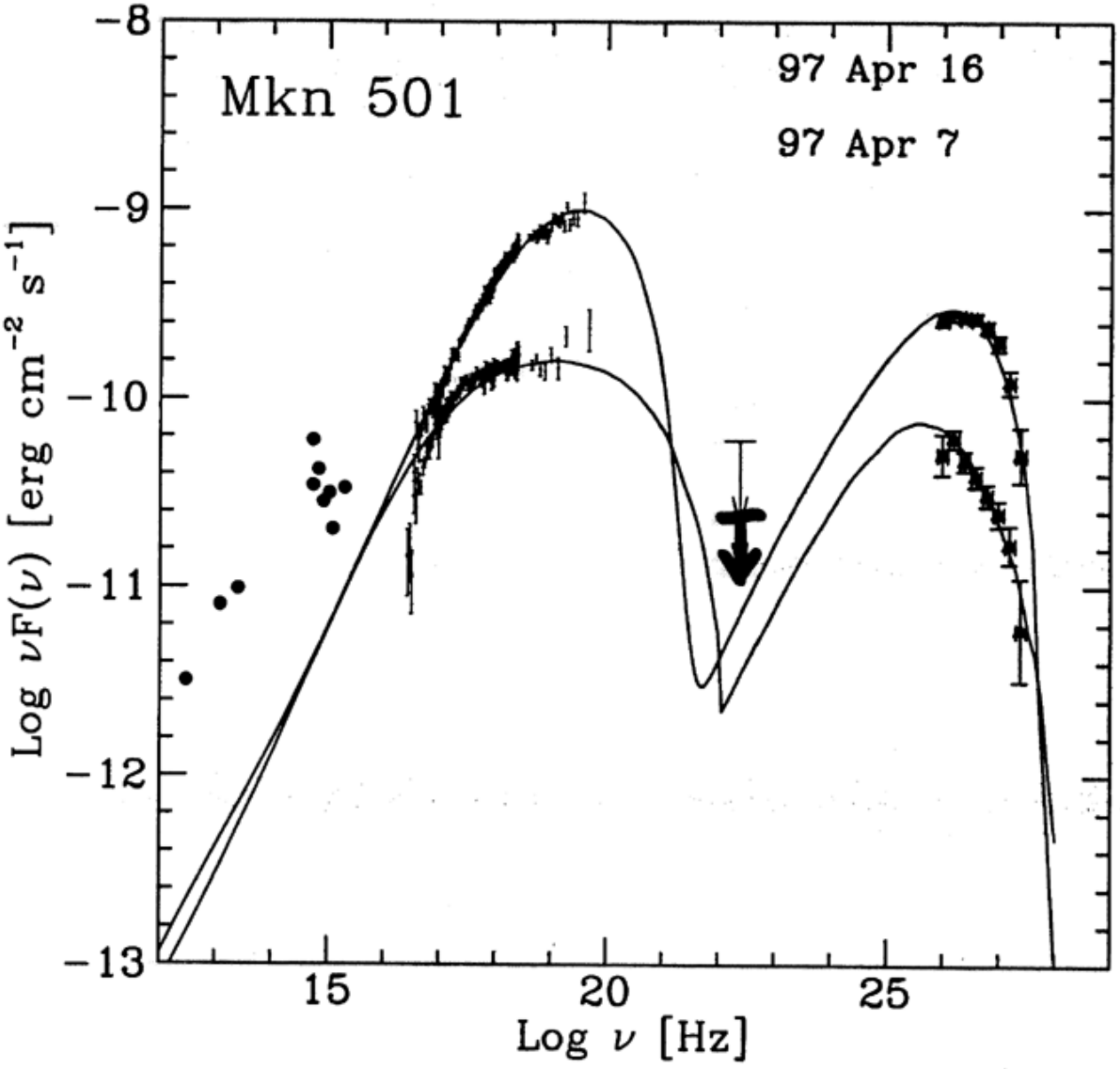
RMS

3C 279 1996 Jan



Wehrle + 1998

RHS



Upper limits are useful!

RMS

λλ coverage for GLAST

- ground based
  - * network
  - * involve amateur astronomers (e.g. BL Lac) and smaller Universities (e.g. Foggy Bottom, Crete, Torino, Perugia ...)

- UV missing!
  XRONOS

- Xrays getting problematic
  RXTE, SAX ?
  AstroE2
  XRONOS

- [illegible] !

RMS

~ 5,000 AGN detected by GLAST

➡ statistical samples
luminosity trends
luminosity function

BUT

need counterpart first!

GLAST error box : ~ 6'

Chandra : ~ 0.16 $\frac{\text{sources}}{\text{arcmin}^2}$

$F_X(2-8\ \text{keV}) \sim 4 \times 10^{-15}$ erg/cm^2/s

S3, 44 ks

(Brandt + 2000)

➡ ~ 6 sources in 6' x 6'

RMS

Need cross-correlations

X rays

optical

radio

dedicated telescopes?

enough workforce

RHS

The unknown:

GLAST potential for new discoveries

AGN

- **Radio galaxies**
  - beaming
  - SSC vs. EC

Cen A — Nolan + 96, Hartman + 99

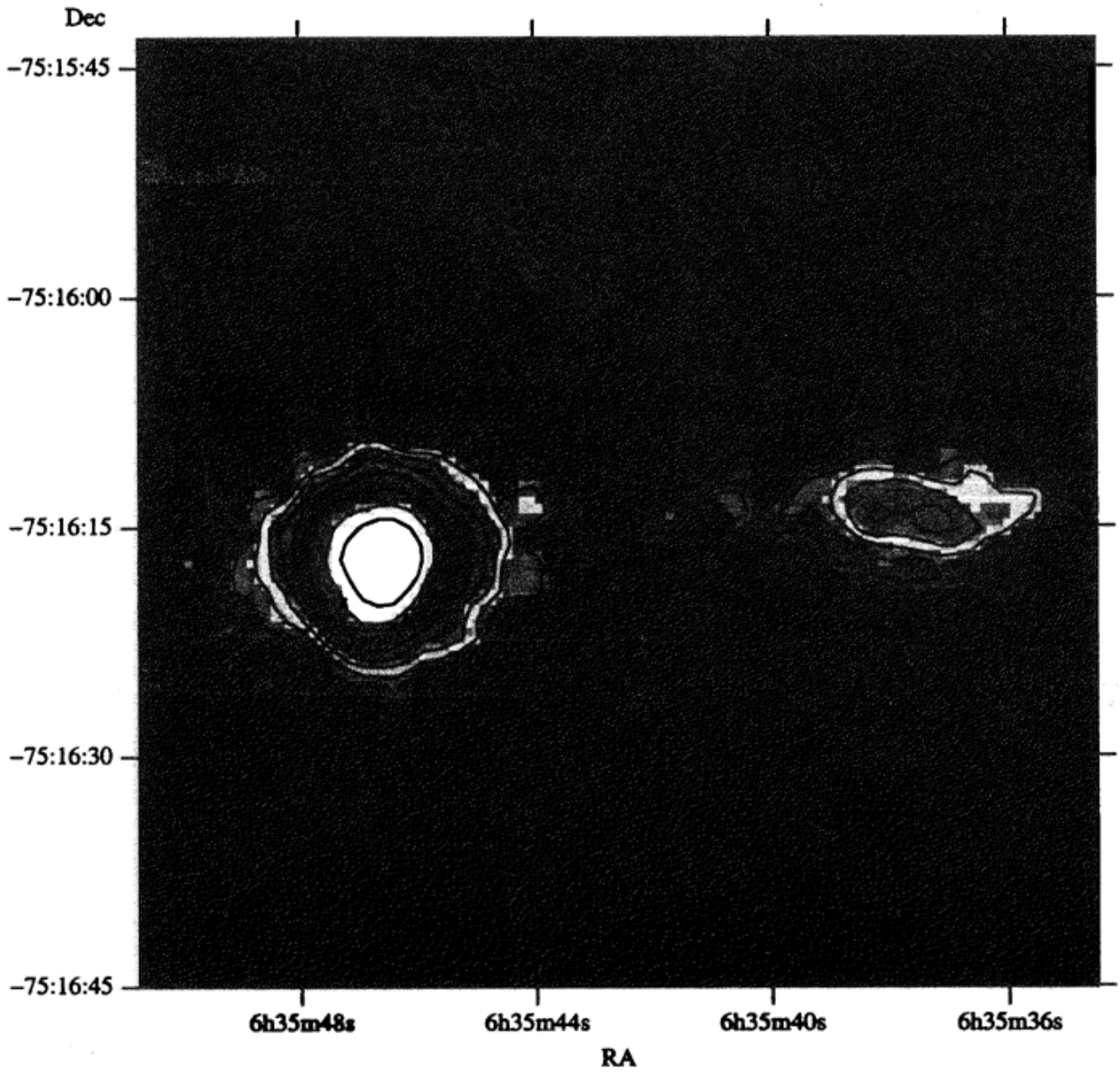
M87, Pic A, ...

Extended jets (e.g. PKS 0637, 3C 273 ...)

RHS

PKS 0637-752 (z = 0.654)

Chandra 100 ks

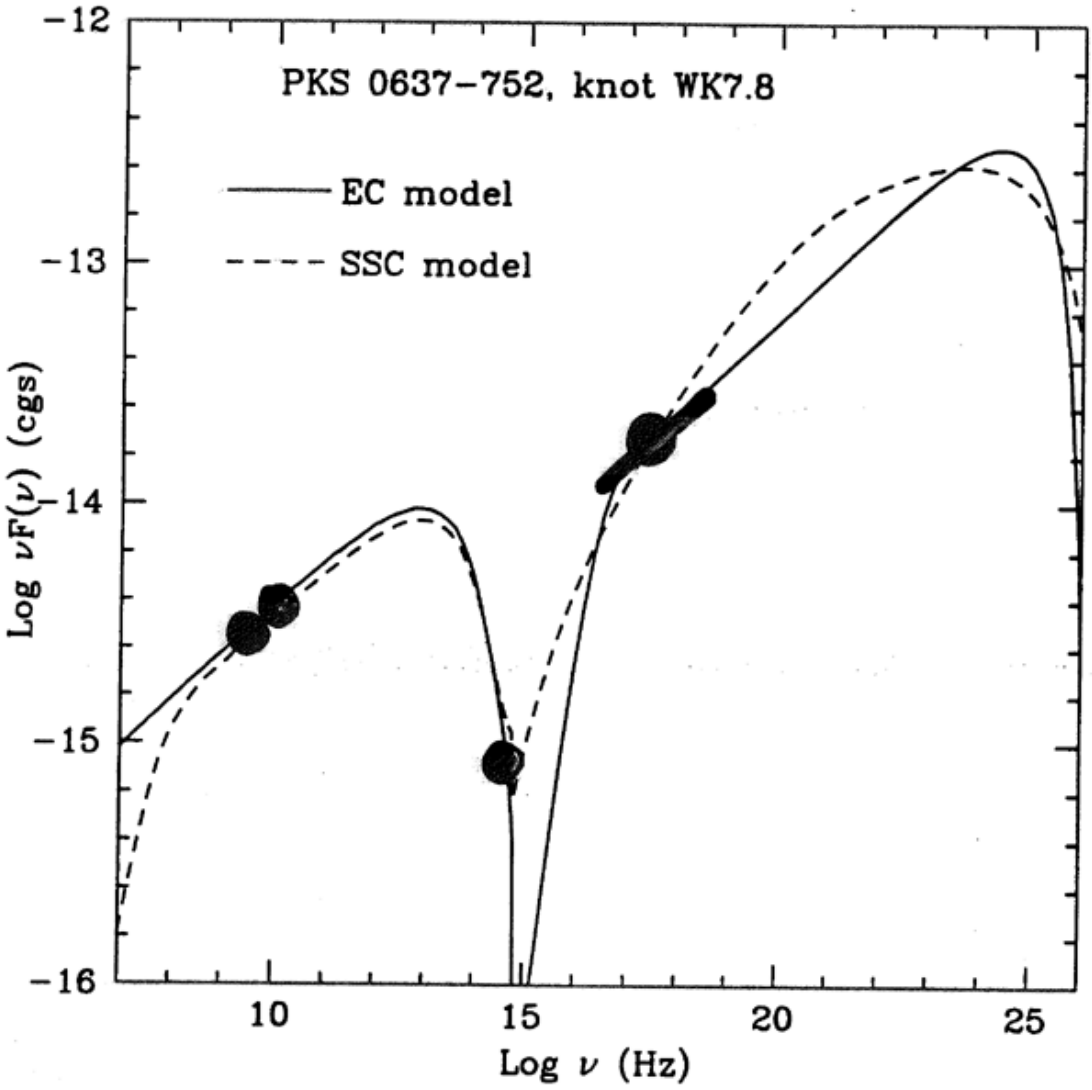


Chartas + 2000

PKS

PKS 0637-752

SED knot WK7.8



Schwartz et al. 2000
Tavecchio et al. 2000

- Seyferts, LLAGNs, LINERS... (RHS)

  Compact radio cores,
  SL, jets

  Falcke + 2000
  Brunthaler + 2000
  Nagar + 2000

origin of radio-loud / radio-quiet AGN ?